

Obscured and powerful AGN and starburst activities at $z \sim 3.5^*$

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ABSTRACT

Aims. Short phases of coeval powerful starburst and AGN activity during the lifetimes of the most massive galaxies are predicted by various models of galaxy formation and evolution. In spite of their recurrence and high luminosity, such events are rarely observed. Finding such systems, understanding their nature, and constraining their number density can provide key constraints to galaxy evolutionary models and insights into the interplay between starburst and AGN activities.

Methods. We report the discovery of two sources at $z = 3.867$ and $z = 3.427$ that exhibit both powerful starburst and AGN activities. They benefit from multi-wavelength data from radio to X rays from the CFHTLS-D1/SWIRE/XMDS surveys. Follow-up optical and near-infrared spectroscopy, and millimeter IRAM/MAMBO observations are also available. We performed a multi-wavelength analysis of their spectral energy distributions with the aim of understanding the origin of their emission and constraining their luminosities. A comparison with other composite systems at similar redshifts from the literature is also presented.

Results. The AGN and starburst bolometric luminosities are $\sim 10^{13} L_\odot$. The AGN emission dominates at X ray, optical, mid-infrared wavelengths, and probably also in the radio. The starburst emission dominates in the far-infrared. The estimated star formation rates range from 500 to 3000 M_\odot/yr . The AGN near-infrared and X ray emissions are heavily obscured in both sources with an estimated dust extinction $A_V \geq 4$, and Compton-thick gas column densities. The two sources are the most obscured and most luminous AGNs detected at millimeter wavelengths currently known.

Conclusions. The sources presented in this work are heavily obscured QSOs, but their properties are not fully explained by the standard AGN unification model. In one source, the ultraviolet and optical spectra suggest the presence of outflowing gas and shocks, and both sources show emission from hot dust, most likely in the vicinity of the nucleus. Evidence of moderate, AGN-driven radio activity is also found in both sources. Based on the estimated stellar and black hole masses, the two sources lie on the local $M_{\text{BH}}-M_{\text{bulge}}$ relation. To remain on this relation as they evolve, their star formation rate has to decrease or stop. Our results support evolutionary models that invoke radio feedback such as the star formation quenching mechanism, and suggest that such a mechanism might play a major role also in powerful AGNs.

Key words. galaxies: active – galaxies: evolution – galaxies: high-redshift – quasars: general – infrared: galaxies

1. Introduction

The similarity between the star formation history and the space density of active galactic nuclei (AGNs) over cosmic time (e.g. Hartwick 2004; Hasinger et al. 2005) and the correlation in

nearby galaxies between black hole (BH) mass and bulge mass (Ferrarese 2002; Gebhardt et al. 2000) indicate that star formation in a galaxy is related to the growth of its BH (but see Shields et al. 2008). Theoretical models (Granato et al. 2001, 2004; Di Matteo et al. 2005; Hopkins et al. 2005a; Springel et al. 2005) explain the link between star formation and AGN activity. They show that gas-rich galaxy mergers are viable precursors to the formation of both massive galaxies and super massive BHs (SMBHs). According to these models, QSOs and (sub)-millimeter galaxies (hereinafter SMGs), the most intense sites of star formation at high- z , represent different stages in an evolutionary sequence, and the obscured growth phase of the BH coincides with the transition from the SMG to the QSO stages (Page et al. 2004; Stevens et al. 2005). QSOs are thus expected to be the end-product of the rapid growth of SMBHs seen soon after a phase of high star formation, as seen in SMGs (Sanders et al. 1988; Granato et al. 2001). Although these models claim to be successful in reproducing and

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explaining several observations (Hopkins et al. 2005b, 2006; Li et al. 2007; Chakrabarti et al. 2007), more observations are necessary to test their predictions, constrain their parameters, and provide a physical base to some of their assumptions (see e.g. Marulli et al. 2008).

Previous studies of SMG populations and high- z starbursts have revealed the presence of obscured AGN activity in ~ 30 – 40% of these systems (e.g. Page et al. 2004; Chapman et al. 2005; Stevens et al. 2005; Alexander et al. 2005c; Yan et al. 2007; Pope et al. 2008), giving support to a link between the two activities. However, it is not clear how star formation and BH growth are linked, what their relative timescales are, and whether they influence each other directly through a feedback mechanism or indirectly, e.g. by consuming the available cold gas. The number of well studied composite systems is still low, and it is often difficult to separate and quantify the contribution from different energy sources (see e.g. the variety of models proposed to explain the properties of F10214+4724; Rowan-Robinson et al. 1993; Teplitz et al. 2006; Efstathiou 2006). In order to quantify their contribution to the bolometric luminosity, the multi-wavelength spectral energy distributions (SEDs) of such systems need to be measured and modeled. Because of the difficulty of obtaining a full multi-wavelength coverage, AGN, starburst galaxies, and composite systems with well sampled SEDs and known spectra are often used to derive correlation between bolometric luminosities and measurements at specific wavelengths or of spectral features (e.g. Sajina et al. 2008; Polletta et al. 2008). In case of powerful starburst galaxies, like SMGs, the AGN is often highly obscured and requires either ultra deep X-ray observations (Alexander et al. 2005a), spectropolarimetric observations (Goodrich et al. 1996), or mid-infrared (MIR) spectroscopic observations (Pope et al. 2008) to be revealed.

In order to investigate the link between star formation and AGN activity, it is thus important to study systems where both activities are taking place, and collect multi-wavelength measurements to constrain the starburst, i.e. the star formation rate (SFR), and AGN, i.e. accretion rate, luminosities, as well as their Eddington ratio, and stellar and gas masses.

Here, we investigate the properties of two rare millimeter (mm) bright obscured QSOs at high- z discovered in a wide multi-wavelength survey. Spectroscopic data as well as the full SED, from X ray to radio wavelengths are available for both objects and are analyzed to constrain the origin of their luminosity and investigate their nature. Throughout this paper, we adopt a flat cosmology with $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.27$ and $\Omega_\Lambda = 0.73$ (Spergel et al. 2003).

2. Target selection and observations

In this work, we analyze the properties of two sources that were selected as g drop-out sources, thus $z \simeq 4$ candidate, in the 0.9 deg^2 CFHTLS survey D1¹ (RA = $02^{\text{h}}26^{\text{m}}$, Dec = $-04^{\circ}30'$) and as very bright MIR sources in the XMM-LSS field of the *Spitzer* Wide-area InfraRed Extragalactic Legacy survey (SWIRE²; Lonsdale et al. 2003). The sources IAU official names are SWIRE2 J022550.67–042142.2 (SW022550 hereinafter), and SWIRE2 J022513.90–043419.9 (SW022513 hereinafter). We applied a modified version of the drop-out selection technique developed by Steidel et al. (1999) for $z > 3$ – 4 Lyman-break sources. In addition to the optical ugr colors and

the source extension, this selection technique takes into account the 3.6 – $4.5 \mu\text{m}$ color (Bergeron et al., in prep.; see also Siana et al. 2008). No other sources with these properties were found in the 0.9 deg^2 CFHTLS-D1/SWIRE field. Because of their large $24 \mu\text{m}$ fluxes ($F_{24 \mu\text{m}} > 1 \text{ mJy}$), exceptionally red infrared (IR) SEDs ($\alpha_{\text{IR}} < -2.5$ where $F_\nu \propto \nu^{\alpha_{\text{IR}}}$ over the observed wavelength range 3 – $24 \mu\text{m}$), and high- z , we selected both sources for observations at 1.2 mm with the Max Planck Millimeter Bolometer (MAMBO) array (Kreysa et al. 1998) at the Institut de Radioastronomie Millimétrique (IRAM) 30 m telescope. Follow up spectroscopic observations were carried out and broad-band photometric data from X ray to radio wavelengths were already available in the field.

2.1. Optical and infrared imaging

Broad-band photometric data at optical and IR wavelengths are available for both sources from various surveys. Optical data in 5 broad-bands, $ugriz$, were provided by the CFHTLS survey D1 (data release T0004³). Near-IR (NIR) data in the J and K bands were obtained from the UKIRT Infrared Deep Sky Survey (UKIDSS⁴ data release 3; Dye et al. 2006; Lawrence et al. 2007). Infrared data in all seven *Spitzer* bands, from IRAC (3.6, 4.5, 5.8, and $8.0 \mu\text{m}$; Fazio et al. 2004) and MIPS (24, 70, and $160 \mu\text{m}$; Rieke et al. 2004), are available from the SWIRE survey. The SWIRE data correspond to the latest internal catalogs that will be released as part of the Data Release 5 (DR5; for details on the data reduction see Surace et al. 2005). The total measured magnitudes and fluxes in each band are reported in Table 1.

2.2. MAMBO observations

Observations at the IRAM 30 m telescope were carried out during March, April and November 2006, using the 117 element version of the MAMBO array operating at a wavelength of 1.2 mm (250 GHz). We used the standard on-off photometry observing mode, chopping between the target and sky at 2 Hz, and nodding the telescope every 10 or 20 s. On-off observations were typically obtained in blocks of 6 scans of 16 or 20 10 s-subscans each, and repeated in later observing nights. The atmospheric transmission was intermediate with $\tau(1.2 \text{ mm}) = 0.1$ – 0.4 . The absolute flux calibration was established by observations of Mars and Uranus, resulting in a flux calibration uncertainty of about 20%.

On average, the noise of the channel used for point-source observations was about 35 – $40 \text{ mJy}/\sqrt{t}/\text{beam}$, where t is the exposure time in seconds, consistent with the MAMBO time estimator for winter conditions. The two sources were observed as part of a larger program targeting more than 100 starburst, AGN, and composite sources selected in the SWIRE fields (Lonsdale et al. 2008; Polletta et al., in prep.; Fiolet et al., in prep.). The resulting 1.2 mm fluxes are $4.7 \pm 0.8 \text{ mJy}$ and $5.5 \pm 0.7 \text{ mJy}$ for SW022550, and SW022513, respectively.

2.3. Optical spectroscopy of SW022550

An optical spectroscopic observation of SW022550 was carried out with the AAOmega system (Sharp et al. 2006) on the 3.9-m Anglo-Australian Telescope, on September 27, 2006, as part of a multi-object spectroscopic program of QSOs in the CFHTLS

¹ <http://www.cfht.hawaii.edu/Science/CFHTLS/>

² <http://swire.ipac.caltech.edu/swire/swire.html>

³ http://terapix.iap.fr/rubrique.php?id_rubrique=241

⁴ <http://www.ukidss.org/>

Table 1. Multi-wavelength data of SW022550 and SW022513.

Mag or flux ^a	Project/Telescope	SW022550	SW022513
$F_{0.5-2\text{ keV}}$	XMDS/XMM	<2.5	1.1 ± 0.8
$F_{2-10\text{ keV}}$	XMDS/XMM	<20	7.8 ± 4.1
m_u	CFHTLS/CFHT	25.211 ± 0.077	>26.00
$m_{g'}$	CFHTLS/CFHT	22.944 ± 0.007	24.432 ± 0.033
m_r	CFHTLS/CFHT	21.648 ± 0.003	23.287 ± 0.014
m_i'	CFHTLS/CFHT	22.949 ± 0.004	22.851 ± 0.012
m_z	CFHTLS/CFHT	22.408 ± 0.021	22.503 ± 0.027
F_J	UKIDSS/UKIRT	3.6 ± 0.4	3.1 ± 0.6
F_K	UKIDSS/UKIRT	7.7 ± 0.6	12.7 ± 0.6
$F_{3.6\text{ }\mu\text{m}}$	SWIRE/Spitzer	15 ± 1	16 ± 1
$F_{4.5\text{ }\mu\text{m}}$	SWIRE/Spitzer	23 ± 1	20 ± 1
$F_{5.8\text{ }\mu\text{m}}$	SWIRE/Spitzer	<58	28 ± 4
$F_{8.0\text{ }\mu\text{m}}$	SWIRE/Spitzer	239 ± 7	180 ± 5
$F_{24\text{ }\mu\text{m}}$	SWIRE/Spitzer	3.32 ± 0.02	2.35 ± 0.02
$F_{70\text{ }\mu\text{m}}$	SWIRE/Spitzer	<24	<24
$F_{160\text{ }\mu\text{m}}$	SWIRE/Spitzer	<126	<126
$F_{1.2\text{ mm}}$	IRAM/MAMBO	4.70 ± 0.77	5.53 ± 0.72
$F_{20\text{ cm}}$	VIMOS/VLA	0.14 ± 0.02	0.35 ± 0.02
$F_{50\text{ cm}}$	VIMOS/GMRT	0.41 ± 0.04	0.62 ± 0.04

^a X ray fluxes are in 10^{-15} erg cm $^{-2}$ s $^{-1}$. Optical magnitudes are total magnitudes in the AB systems in the $ugriz$ filters from the CFHTLS survey D1. UKIDSS J and K fluxes, and IRAC (3.6–8.9 μm) fluxes are in μJy . MIPS (24–160 μm), MAMBO (1.2 mm) and radio (20 cm) fluxes are in mJy. X-ray upper limits correspond to 3σ , optical upper limits correspond to 80% completeness, IR upper limits correspond to 5σ . The radio data are from the VIMOS VLA survey (Bondi et al. 2003, 2007).

(Petitjean et al., in prep.; Yu et al. in prep.). The 580V and 385R gratings were used in the blue and red arms of the spectrograph, respectively, yielding $R \sim 1300$. The total integration time was 3 h, divided into six 30 min exposures.

Data reduction was performed using the AAOmega data reduction pipeline software DRCONTROL. The two dimensional images were flatfielded, and the spectrum was extracted (using a Gaussian profile extraction), wavelength calibrated and combined within DRCONTROL. The final spectrum was flux calibrated using the broad-band photometric measurements in the g and r bands after deconvolving the spectrum with the filter bandpasses. This calibration assumes that the source optical emission has not significantly varied between the spectroscopic and photometric observations.

The optical spectrum of SW022550 shows several strong emission lines, e.g. Ly α , N V $\lambda 1240$, C IV $\lambda 1549$, the O VI $\lambda 1035$ doublet (see Fig. 1). The continuum emission is weak ($<2\sigma$), and a drop in the continuum emission is visible at rest-frame wavelengths below 1216 \AA due to intergalactic medium (IGM) absorption. The bottom panel of Figure 1 shows the decrease in the continuum emission at rest-frame wavelengths $\lambda < 1216\text{ }\text{\AA}$. Absorption features due to the Lyman α forest are also visible. Since the line spread function (LSF) derived from the lamp spectra is consistent with a Gaussian, we used such a profile to fit the emission lines. From the fits we measured the total line flux, the central wavelength, the $FWHM$ in \AA , and the flux of the local linear continuum. The $FWHM$ in km s^{-1} and the rest-frame equivalent width (W_λ) in \AA are then derived from these parameters after correcting for instrumental resolution. The instrumental $FWHM$ is 3 \AA at $\lambda < 5500\text{ }\text{\AA}$, and 5 \AA at $\lambda > 5500\text{ }\text{\AA}$. The line flux, central wavelength, corresponding redshift, rest-frame $FWHM$ and rest-frame W_λ of the main visible lines are reported in Table 2. Uncertainties are determined

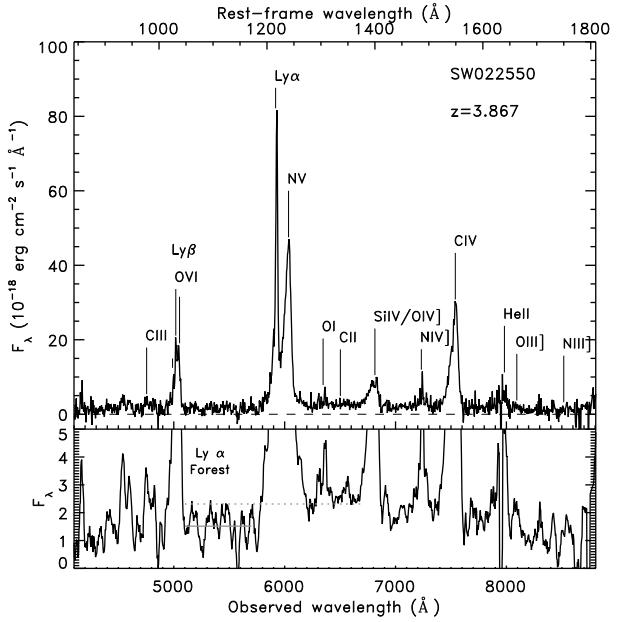


Fig. 1. AAOmega optical spectrum of SW022550. The expected emission lines are labeled. The horizontal dashed line corresponds to a null flux. The *bottom panel* shows the same spectrum but smoothed with a kernel of 7 pixels to better illustrate the continuum emission. The top dotted grey line represents the mean continuum level at rest-frame $\lambda = 1275\text{--}1380\text{ }\text{\AA}$, and the bottom solid grey curve the mean continuum level at $\lambda = 1050\text{--}1170\text{ }\text{\AA}$.

using a Montecarlo method and assuming a Gaussian noise given by the rms associated with the continuum where the fit is performed. A broad component, in addition to a narrow one, is measured in correspondence of the Ly α , N V $\lambda 1240$, Si IV $\lambda 1397$, and C IV $\lambda 1549$ lines. All these lines show an asymmetric component, similar to a blue wing. To fit such an asymmetric profile, in some cases, e.g. Ly α , and N V $\lambda 1240$, we fit multiple Gaussian components. Based on the central wavelength of the optical lines listed in Table 2, we estimate a redshifts of 3.867 ± 0.009 .

The emission lines are dominated by their narrow ($FWHM \sim 1000\text{ km s}^{-1}$) components with $FWHMs$ typical of type 2 QSOs ($\leq 1500\text{ km s}^{-1}$; see e.g. Baldwin et al. 1988; Norman et al. 2002; Mainieri et al. 2005). Broad components are also present with $FWHM \sim 2000\text{--}9000\text{ km s}^{-1}$ as observed in type 1 QSOs ($\geq 2000\text{ km s}^{-1}$; see e.g. Richards et al. 2004). The equivalent widths of the emission lines in the optical spectrum of SW022550 are larger than observed in type 1 QSOs, and more similar to those observed in other type 2 QSOs (see e.g. SMM 02399–0136, CDFS-263, or CXO 52; Ivison et al. 1998; Mainieri et al. 2005; Stern et al. 2002). In Fig. 2, we compare the optical spectrum of SW022550 with a composite spectrum of a large sample of type 1 QSOs and with the spectra of four type 2 QSOs at high- z from the literature. We show the spectra normalized at the Ly α peak on the top panel, and at the C IV $\lambda 1549$ peak in the bottom panel. The composite type 1 QSO spectrum corresponds to the median composite spectrum of 2204 QSO spectra from SDSS (Vanden Berk et al. 2001). The type 2 QSOs have been selected because of the availability of their optical spectra from the literature and high- z . They are CDFS-202, and CDFS-263 from CDFS (Szokoly et al. 2004), and SW104406 and SW104409 from the SWIRE/Chandra survey (Polletta et al. 2006). This comparison shows the wide range of line ratios that can be observed in type 2 QSOs. The main differences reside in the strength of the N V $\lambda 1240$, and C IV $\lambda 1549$ emission

Table 2. Emission-line measurements.

Line	λ_{obs} (Å)	z	$W_{\lambda, \text{rest}}$ (Å)	Flux (10^{-16} erg cm $^{-2}$ s $^{-1}$)	$FWHM^a$ (km s $^{-1}$)	Comments
SW022550						
Ly β	4997.4 ± 0.8	3.876	41 ± 3	1.35 ± 0.11	935 ± 180	Narrow
[O VI] $\lambda\lambda 1031, 1038$	5018.2 ± 0.6	3.867	68 ± 3	2.25 ± 0.10	777 ± 81	...
	5042.9 ± 0.8	3.858	152 ± 4	5.03 ± 0.14	1820 ± 114	...
Ly α	5928 ± 0.1	3.875	322 ± 5	13.52 ± 0.19	1041 ± 13	Narrow
N V $\lambda 1240$	6022 ± 0.7	3.856	311 ± 9	13.19 ± 0.39	2531 ± 99	Broad
	6044 ± 0.7	3.874	146 ± 5	6.19 ± 0.20	1354 ± 74	Narrow
Ly α	5920 ± 1.8	3.868	164 ± 15	6.88 ± 0.65	4060 ± 157	Broad
	5985 ± 2.2	...	496 ± 23	20.98 ± 0.96	8968 ± 343	Broad
Si IV $\lambda 1397$	6778 ± 4.3	3.852	833 ± 37	7.38 ± 0.33	4581 ± 318	...
	6820 ± 2.4	3.861	281 ± 22	2.49 ± 0.20	2375 ± 172	...
N IV] $\lambda 1486$	7242 ± 2.4	3.873	41 ± 4	1.95 ± 0.18	1987 ± 239	...
C IV $\lambda 1549$	7527 ± 1.5	3.859	328 ± 8	16.13 ± 0.38	3690 ± 101	Broad
	7542 ± 0.8	3.869	91 ± 4	4.47 ± 0.18	1235 ± 88	Narrow
He II $\lambda 1640$	7970 ± 15	3.860	84 ± 8	2.75 ± 0.26	3098 ± 768	...
H β	23 560	3.847	$>37^b$	1.4 ± 0.5	$<510^c$	ISAAC
[O III] $\lambda 4959$	24 060	3.852	$>239^b$	9.1 ± 0.4	$<498^c$	ISAAC
SW022513						
H β	$21\,532 \pm 5$	3.429	$>120^b$	<2.5	$<557^c$	ISAAC
[O III] $\lambda 4959$	$21\,954 \pm 2$	3.427	$>297^b$	6.1 ± 0.5	288 ± 58	ISAAC
[O III] $\lambda 5007$	$22\,160 \pm 1$	3.426	$>829^b$	16.9 ± 0.6	660 ± 34	ISAAC

All measurements are based on single Gaussian fits to the emission lines assuming a flat (in F_{λ}) continuum. ^a Corrected for instrumental resolution.

^b Since the continuum is not detected we estimate an upper limit to the equivalent width, W_{λ} , assuming a 3σ upper limit to the continuum. The continuum is $<3.78 \times 10^{-18}$ erg cm $^{-2}$ s $^{-1}$ Å $^{-1}$ for SW022550, and $<2.03 \times 10^{-18}$ erg cm $^{-2}$ s $^{-1}$ Å $^{-1}$ for SW022513. ^c The $FWHM$ is fixed to the spectral resolution because it cannot be correctly determined due to the low signal-to-noise.

lines, and the blue wings observed in the main emission lines. We will discuss below two possible explanations for the strength of the N V $\lambda 1240$, and C IV $\lambda 1549$ lines, i.e. high metallicities and shocks. Interestingly, there is another object in the type 2 QSO sample from the literature with similar asymmetric lines, i.e. SW104409. This source is also characterized by a similar optical-IR SED to SW022550, with an unusual optical blue continuum with strong AGN emission lines, and red optical-MIR colors. A possible explanation for the presence of blue wings is extinction due to dust mixed with infalling or outflowing ionized gas (Osterbrock 1989). Since the SED in both cases indicate that the AGN is obscured, it is also plausible that the blue continuum and the broad blue wings are due to scattered light (see e.g. Zakamska et al. 2006).

SW022550 is characterized by large N V $\lambda 1240$ /C IV $\lambda 1549$ and N V $\lambda 1240$ /He II $\lambda 1640$ flux ratios. These flux ratios are often used to estimate the gas metallicities in QSOs, from either the broad or the narrow line region (BLR and NLR; see Nagao et al. 2006, and references therein). High values usually imply high metallicities (see e.g. Hamann & Ferland 1993; Vernet et al. 2001; Norman et al. 2002). Based on the photoionization model in Hamann & Ferland (1993), we derive a metallicity $Z = 4 Z_{\odot}$. However, not all the emission line ratios of the spectrum of SW022550 are explained by NLR photoionization models (see e.g. Nagao et al. 2006; Groves et al. 2004). A possible cause of this discrepancy is the presence of broad components. It is also possible that the emission lines are produced by shocks in addition to, or rather than by photoionization. This hypothesis is supported by the strength of some lines, like C IV $\lambda 1549$, O VI $\lambda 1035$, and N V $\lambda 1240$. For example, C IV $\lambda 1549$ is predominantly produced in the cooling zone of shocks and it is expected to be stronger in shocks than in photoionization processes. Indeed, the C IV $\lambda 1549$ /He II $\lambda 1640$, N V $\lambda 1240$ /He II $\lambda 1640$, and N V $\lambda 1240$ /C IV $\lambda 1549$ flux

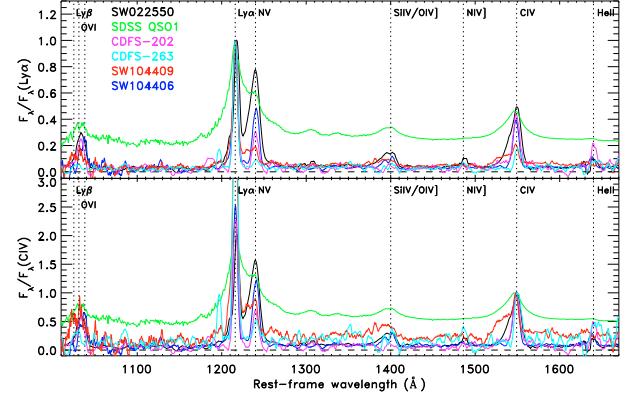


Fig. 2. Top panel: optical spectra normalized at the flux at 1216 Å in the rest-frame of SW022550 (black thick line), of the type 2 QSOs CDFs-263 (cyan line) and CDFs-202 (magenta line) from Szokoly et al. (2004), and SW104406 (blue line), and SW104409 (red line) from Polletta et al. (2006), and medium composite spectrum of SDSS type 1 QSOs (green line; Vanden Berk et al. 2001). Bottom panel: as top panel, but normalized at the flux at 1549 Å rest-frame. The main emission lines are labeled and marked with dotted vertical lines. The horizontal dashed line corresponds to a null flux. All spectra have been smoothed for clarity.

ratios like those observed in SW022550, are higher than predicted by photoionization, and can be explained by shocks (Allen et al. 2008), or by the coexistence of shocks and AGN photoionization, without requiring largely super-solar metallicities (Moy & Rocca-Volmerange 2002). Also the presence of O VI $\lambda 1035$ can be considered as evidence of shocks, although other explanations are also possible. To confirm whether SW022550 is characterized by high metallicities and better constrain its value, or whether shocks are in part responsible for the

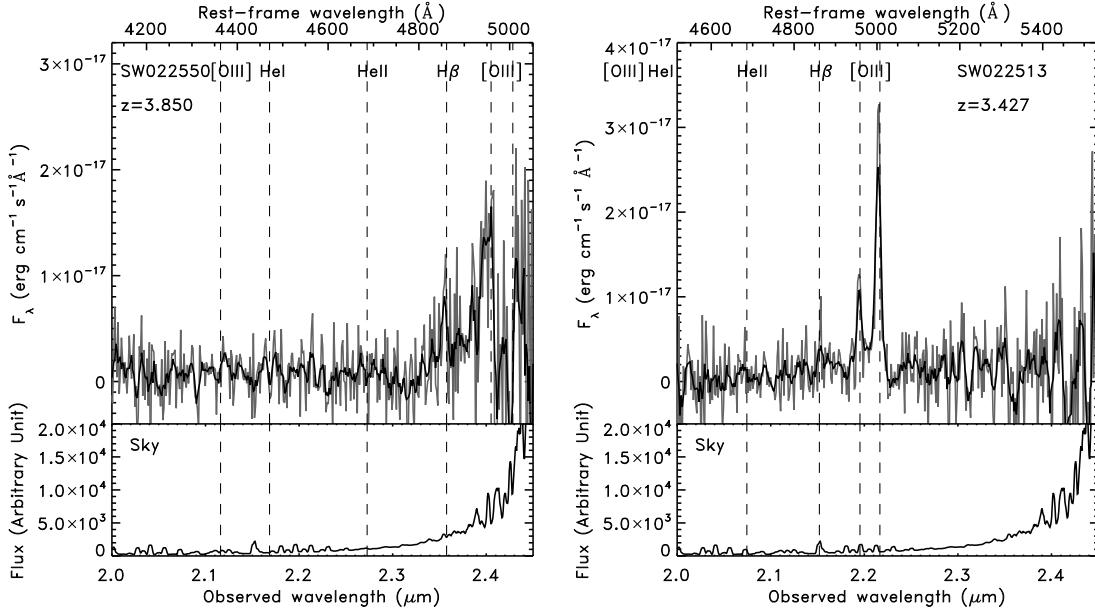


Fig. 3. The top panels show the ISAAC SK spectra of SW022550 (left panel) and SW022513 (right panel). The grey curves correspond to the spectrum smoothed with a kernel of 2 pixels, and the black curves to the smoothed spectrum with a kernel of 10 pixels. The bottom panels show the SK band sky spectrum. The location of the main expected lines are labeled.

observed emission lines, it would be necessary to measure other emission lines, especially at rest-frame optical wavelengths.

In summary, the optical spectrum of SW022550 shows spectral features that are also observed in other obscured (type 2) AGNs at high- z , but are not explained by the standard AGN unification model (Antonucci 1993). In particular, a high metallicities or the presence of shocks and a significant scattered fraction at rest-frame UV wavelengths are necessary to fully explain the observed spectrum. Note that these hypothesis can not be fully tested with the available data and other scenarios are not ruled out.

2.4. Near-infrared spectroscopy

Both targets were observed with the ISAAC instrument (Moorwood et al. 1998) on Antu (VLT-UT1) in low-resolution (LR) mode, using the 1024×1024 Hawaii Rockwell array of the Short Wavelength arm. Observations were carried out with a $1''$ slit with two grating blaze angles, approximately covering the SH ($\lambda \sim 1.42$ – $1.82 \mu\text{m}$), and SK ($\lambda \sim 1.82$ – $2.5 \mu\text{m}$) bands. The nominal resolutions were 500, and 450 for the two cases, respectively. The observations of the two targets were carried out in visitor mode on September 13th–14th, 2007. Total integration times were 81 min and 66 min in the SH band, and 84 min and 42 min in the SK band on SW022550, and SW022513, respectively.

Each observation was splitted in sets of 180 s exposures, adopting a nodding and jittering pattern for optimal background subtraction. The nod throw was set to 60 – $80''$, in order to avoid conflict with the reference star; the jitter width was $10''$.

Data reduction was carried out in the standard manner, using the IRAF environment⁵, including ABBA sky subtraction. Wavelength calibration was based on OH sky lines. Atmospheric

extinction correction and flux calibration were obtained with standard telluric stars, carefully selected from the ESO database in order to have known NIR magnitudes. For this purpose, several B-type telluric standards were observed each night (before and after a target was observed) at air masses within 0.1–0.2 of the air mass of the target observations. The rebinned ISAAC SK-band spectra of the two sources are shown in Fig. 3, and discussed below.

2.4.1. ISAAC SK Spectrum of SW022550

Neither continuum nor lines are detected in the SH-band spectrum of SW022550. Two emission lines are detected in the SK-band spectrum at $\lambda 23\,560 \text{ \AA}$, and $24\,060 \text{ \AA}$ yielding a $z \sim 3.85$ if interpreted as $\text{H}\beta$, and $[\text{O III}] \lambda 4959$, respectively. The continuum in the SK-band is not detected and it is estimated to be $<3.8 \times 10^{-18} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ (corresponding to 3σ). The identification of these lines is tentative due to the high sky level at these wavelengths and the uncertain wavelength calibration at $\lambda \geq 2.4 \mu\text{m}$. The $[\text{O III}] \lambda 4959$ line shows a blue tail probably due to a poor sky subtraction. At the expected location of the $[\text{O III}] \lambda 5007$ line, the spectrum is negative, suggesting that the sky subtraction might have been too high at these wavelengths. The $\text{H}\beta/\text{Ly}\alpha$ flux ratio is 0.03 ± 0.1 , which is consistent with the range of values observed in type 2 AGNs, i.e. from 0.014 to 0.03, while type 1 AGNs are usually characterized by higher values (0.07–0.125). Note that because of the poor sky subtraction where the lines are observed (see bottom panel of Fig. 3), it is possible that a broad $\text{H}\beta$ component is also present. Assuming a $FWHM$ of 5000 km s^{-1} for a broad $\text{H}\beta$ we estimate an upper limit to the flux $<16 \times 10^{-16} \text{ erg cm}^{-2} \text{ s}^{-1}$. This value is significantly higher than the predicted flux based on the measured $\text{Ly}\alpha$ flux and the $\text{H}\beta/\text{Ly}\alpha$ flux ratio observed in type 1 AGNs. Thus, we cannot rule out the presence of a broad $\text{H}\beta$ line in the spectrum of SW022550. The $[\text{O III}] \lambda 4959/\text{H}\beta$ flux ratio is 6.3 ± 2.1 . This is also closer to the ratios observed in type 2 AGNs, ~ 2.3 , than to those observed in type 1 AGNs, ~ 0.04 . The upper limit to the $\text{H}\beta$ flux yields a $[\text{O III}] \lambda 4959/\text{H}\beta$ flux ratio >0.56 , still

⁵ The package IRAF is distributed by the National Optical Astronomy Observatory which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

higher than in type 1 QSOs. In summary, in spite of its low signal-to-noise, we conclude that the (rest-frame) optical spectrum of SW022550 is more similar to the spectra of type 2 AGNs, than to those of type 1 AGNs. The ISAAC spectrum shows another interesting property. The $H\beta$, and $[\text{O III}] \lambda 4959$ lines are slightly blueshifted with respect to the optical emission lines. Assuming that such a difference is real and that the source is at $z = 3.867$ as derived from the optical spectrum, the blueshifted lines could be explained by outflowing gas with a velocity $\sim 510 \text{ km s}^{-1}$ (see similar cases in [Sajina et al. 2008](#)). Such a velocity is consistent with those assumed in shocks models ([Allen et al. 2008](#)).

2.4.2. ISAAC SK spectrum of SW022513

Neither continuum nor lines are detected in the SH-band spectrum of SW022513. In the SK-band spectrum, two emission lines at $\lambda_{\text{obs}} = 21954 \text{ \AA}$, and 22160 \AA are well detected, while the continuum is not detected. We thus estimate an upper limit to the continuum $< 2.03 \times 10^{-18} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$ (corresponding to 3σ), and to the emission lines equivalent widths. The two lines are interpreted as $[\text{O III}] \lambda\lambda 4959, 5007$, constraining the redshift to be $z = 3.427 \pm 0.001$. A weak emission feature is observed at the expected wavelength corresponding to the $H\beta$ line. Since the line is affected by a bright sky line which falls exactly at its expected location, and it is narrower than the spectral resolution, we cannot constrain it. The $[\text{O III}] \lambda 4959/H\beta$ and $[\text{O III}] \lambda 5007/H\beta$ flux ratios are, respectively, > 0.97 and > 2.48 . These values are much higher than those observed in type 1 AGNs (i.e. 0.04 and 0.15) and more similar to those observed in type 2 AGNs (2.3 and 5.5) and expected from photoionization models of the narrow line region (NLR; [Osterbrock 1989](#); [Groves et al. 2004](#)). Thus, it is quite likely that such a feature is the $H\beta$ line. In summary, the ISAAC spectrum of SW022513 shows narrow emission lines with flux ratios consistent with those of type 2 AGNs.

The lines fluxes, W_{λ} , and $FWHMs$ of both spectra are reported in Table 2. The SH-band observations cover the rest-frame wavelength range ~ 2900 – 3700 \AA for SW022550, and ~ 3200 – 4100 \AA for SW022513. At these wavelengths, the only feature that might be detected is the $[\text{O II}] \lambda 3727$ emission line. The $[\text{O II}] \lambda 3727$ line is typically associated with star formation activity ([Gallagher et al. 1989](#)), but also with AGN activity. Since our sources are not detected in the SH-band, and an atmospheric feature is present at the expected $[\text{O II}] \lambda 3727$ location in the two sources, we cannot set any constraints on their $[\text{O II}] \lambda 3727$ emission.

From the spectra we estimate that the upper limits to the NIR continuum in SW022513 and SW022550 are $< 33 \mu\text{Jy}$, and $< 61 \mu\text{Jy}$. These values are about 3 and 8 times higher than the K -band measured fluxes. Thus, the spectroscopic measurements are consistent with the broad band photometric data.

2.5. Radio data

Both targets are detected in the VIMOS VLA survey at 1.4 GHz, and at 610 MHz ([Bondi et al. 2003](#); [Ciliegi et al. 2005](#); [Bondi et al. 2007](#)). The measured fluxes are reported in Table 1. In the VLA 1.4 GHz observations, SW022513 appears extended with a major axis of $4.1''$ and a minor axis of $1.8''$, corresponding to a projected linear size of $4.2 \times 2.1 \text{ kpc}$, and a position angle (PA) of 161.3° ([Bondi et al. 2003](#)). In the other radio observations the sources are unresolved. Assuming a power-law model $F_{\nu} \propto \nu^{\alpha}$

for the radio emission, the radio spectral indeces α_r are -1.3 for SW022550, and -0.7 for SW022513. Because of its radio spectral index, SW022550 can be considered an ultra-steep spectrum source (USS, i.e. $\alpha_r \leq -1$; [Roettgering et al. 1994](#), and references therein). Such a steep spectrum implies an AGN origin of the observed emission. In case of SW022513, the spectral index is consistent with what is observed in star forming galaxies and radio-quiet AGNs ([Ciliegi et al. 2003](#)), however the extended emission and the large radio flux strongly suggest also an AGN origin of its radio emission. A more detailed discussion of the radio properties of our sources is presented in Sect. 3.3.

2.6. X ray data

The two targets are also covered by 20ks observations with *XMM-Newton* from the *XMM-Newton* Medium Deep Survey (XMDS; [Chiappetti et al. 2005](#); [Tajer et al. 2007](#)). Only EPIC/MOS data are available as both sources fall in a bad column of the EPIC/pn detector. SW022513 is at $4.5'$ from the aim-point and marginally detected with 29 ± 12 total counts (12 ± 8 at 0.5 – 2 keV , and 17 ± 9 at 2 – 10 keV). The hardness ratio is $HR = 0.18^{+0.7}_{-0.3}$, which corresponds to an effective gas column density $N_{\text{H}} \simeq (1^{+4}_{-0.5}) \times 10^{24} \text{ cm}^{-2}$ assuming a power-law model with photon index $\Gamma = 1.7$, and Galactic and intrinsic photoelectric absorption with $N_{\text{H}}^{\text{Gal}} = 2.61 \times 10^{20} \text{ cm}^{-2}$ ([Dickey & Lockman 1990](#)). More details on the X ray data and derived quantities can be found in [Chiappetti et al. \(2005\)](#), and [Tajer et al. \(2007\)](#). SW022550 is at $> 11'$ from the aim-point and is not detected, thus we can only give an upper limit to its X-ray flux. The X-ray fluxes of both sources are reported in Table 1.

Based on these X ray measurements, the estimated broad and hard X ray luminosities of SW022513 are $\sim 6 \times 10^{44} \text{ erg s}^{-1}$, and $5 \times 10^{44} \text{ erg s}^{-1}$, respectively. The absorption-corrected broad and hard X ray luminosities are $11 \times 10^{44} \text{ erg s}^{-1}$, and $6 \times 10^{44} \text{ erg s}^{-1}$. For SW022550, we estimate an upper limit to the absorbed luminosities in the broad and hard bands of $2.0 \times 10^{45} \text{ erg s}^{-1}$, and $1.8 \times 10^{45} \text{ erg s}^{-1}$, respectively. Based on the estimated X ray luminosity of SW022513 and the relationship between the $[\text{O III}] \lambda 5007$ and the hard X ray luminosity in type 2 AGNs ([Mulchaey et al. 1994](#); [Netzer et al. 2006](#)), we expect an $[\text{O III}] \lambda 5007$ luminosity of 3 – $9 \times 10^{42} \text{ erg s}^{-1}$, instead of $2 \times 10^{44} \text{ erg s}^{-1}$ as observed (see Table 3). Such a discrepancy might be due to the non validity of the X ray/ $[\text{O III}] \lambda 5007$ relationship for our sources, to an underestimation by a factor of 100 of the X ray luminosity, or to the contamination from star formation to the $[\text{O III}] \lambda 5007$ flux. Probably, all these factors play a role in explaining this discrepancy, but the only one that could explain such a large difference is the underestimation of the intrinsic X ray luminosity. A higher X ray luminosity would also be required to obtain a MIR/X-ray luminosity ratio consistent with those observed in other AGNs (see Sect. 3.1). Based on these considerations, it is quite probable that SW022513 is a Compton-thick QSO. Since SW022550 is even fainter in the X-rays, and more luminous in $[\text{O III}] \lambda 5007$ emission and in MIR luminosity, it is also quite plausible that SW022550 is a Compton-thick QSO.

The full radio-Xray SEDs of the two targets are shown in Fig. 5, and all the available fluxes are listed in Table 1.

3. Multi-wavelength spectral energy distributions

The SEDs of SW022550, and SW022513 are shown in Figs. 4 and 5. Both sources are characterized by large

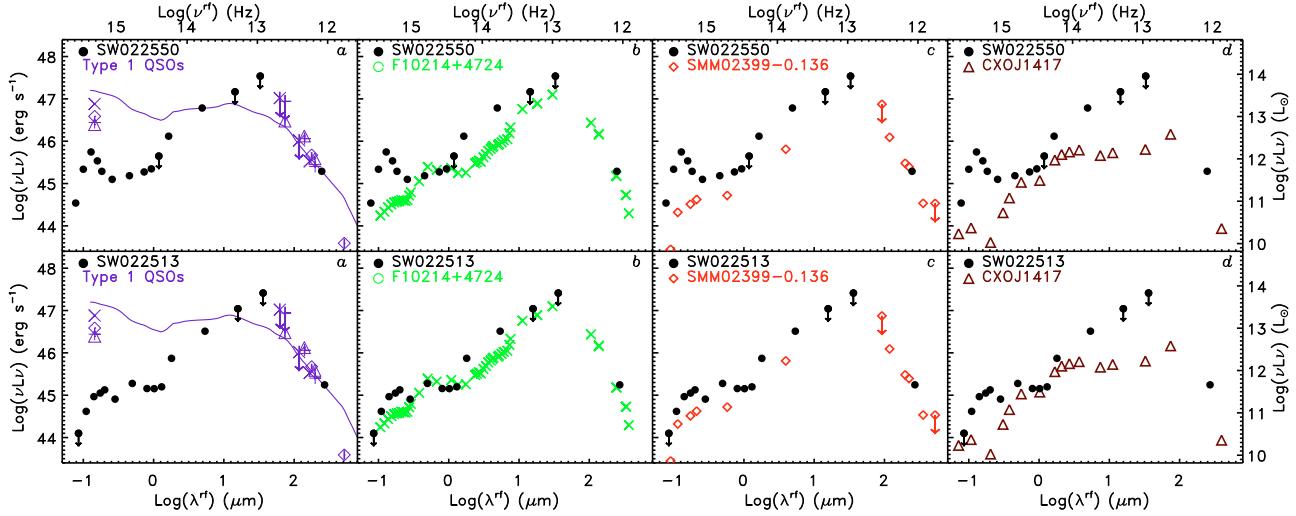


Fig. 4. Rest-frame UV-mm SEDs of the selected targets (black full circles), SW022550 (top panels), and SW022513 (bottom panels), compared with the SEDs of other mm-detected AGNs, **a**) a median type 1 QSO template normalized at the mm luminosity (Elvis et al. 1994), and 4 mm-detected type 1 QSOs with FIR excess (Wang et al. 2007, purple symbols), J033829.31+002156.3 (purple plus signs), J075618.14+410408.6 (purple triangles), J092721.82+200123.7 (purple diamonds), and J104845.05+463718.3 (purple crosses), **b**) F10214+4724 (green crosses; Rowan-Robinson et al. 1993; Teplitz et al. 2006), **c**) SMM02399-0136 (red diamonds; Ivison et al. 1998), and **d**) CXOJ1417 (brown triangles; Le Floc'h et al. 2007).

Table 3. SW022550 and SW022513 luminosities

Parameter ^a	Method ^b	SW022550	SW022513
Log($L(\text{IR})$)	SK07 Starburst	12.5–13.3	12.5–13.2
Log($L_{1.4\text{GHz}}$)	PL	25.50	25.38
Log($L(H)$)	SK07 Starburst	11.2–11.9	11.3–11.6
Log(M_*)	$L(H)$ (S07)	11.2–11.9	11.3–11.6
Log($L(6\text{ μm})$)	H06 Torus	13.1–13.2	12.9–12.9
Log($L_{\text{bol}}^{\text{SB}}$)	SK06 Starburst	12.6–13.4	12.4–13.2
Log($L_{\text{bol}}^{\text{AGN}}$)	H06 Torus	13.4	13.1

^a $L(\text{IR})$ is the 8–1000 μm luminosity in L_\odot derived by integrating the model under Method. $L_{1.4\text{GHz}}$ in the monochromatic radio luminosity at 1.4 GHz in the rest-frame in W Hz^{-1} . $L(H)$ is the monochromatic luminosity in L_\odot at 1.6 μm in the rest-frame derived from the starburst models. M_* is the stellar mass derived from $L(H)$ assuming the average $L(H)/M_*$ ratio derived from the high- z radio galaxy sample in Seymour et al. (2007). $L(6\text{ μm})$ is the rest-frame luminosity at 6 μm in L_\odot derived from the torus models. $L_{\text{bol}}^{\text{SB}}$ is the 0.1–1000 μm luminosity derived from the starburst models in L_\odot . $L_{\text{bol}}^{\text{AGN}}$ is the 0.1–1000 μm luminosity derived from the torus models in L_\odot .

^b The SK07 starburst models are from the library in Siebenmorgen & Krügel (2007). The H06 torus models are from Hönig et al. (2006). PL stands for power-law model with spectral index α_r , $F_\nu \propto \nu^{\alpha_r}$, where $\alpha_r = -1.3$ for SW022550, and -0.7 for SW022513.

MIR fluxes ($F_{24\text{ μm}} \simeq 2\text{--}3 \text{ mJy}$), show red power-laws in the NIR (1–5 μm in the rest-frame; $\alpha_{\text{IR}} < -2.5$, where $F_\nu \propto \nu^{\alpha_{\text{IR}}}$), and large IR/optical flux ratios ($\text{Log}(\nu F_{24\text{ μm}}/\nu F_z) = 1.5$). All these properties are signatures typical of obscured QSOs (Weedman et al. 2006; Polletta et al. 2008).

3.1. Comparison with other high- z QSOs

We compare the UV-mm SEDs of SW022550, and SW022513 with those of well studied mm-detected non radio-loud AGNs at high- z from the literature in Fig. 4. The vast majority of mm-detected QSO at high- z are unobscured or type 1 QSOs with MIR SEDs consistent with the median template of unobscured QSO (e.g. Hines et al. 2006). We represent their SEDs

in panel **a** of Fig. 4 with the median QSO template by Elvis et al. (1994). The majority of these type 1 QSOs with strong mm fluxes, i.e. $\gtrsim 3 \text{ mJy}$, as our sources, show a significant excess in the far-IR (FIR) with respect to the optical emission predicted by the median QSO template (Wang et al. 2007; Hao et al. 2008). The SEDs of 4 QSOs with a FIR excess are also shown in panel **a** of Fig. 4. Our sources show similar mm luminosities, but 10–100 times lower optical fluxes, consistent with being obscured. In Fig. 4, we also compare our sources SEDs with those of 3 mm-detected obscured QSOs with available MIR data, (1) the lensed system F10214+4724, a Seyfert 1 galaxy at $z = 2.86$, but that is also heavily absorbed in the X rays (panel **b**; Ivison et al. 1998; Alexander et al. 2005b); (2) SMM02399-0136, a type 2 QSO at $z = 2.803$ (panel **c**; Ivison et al. 1998); and (3) CXO GWS J141741.9+522823 (CXOJ1417 hereinafter) at $z = 1.15$ (panel **d**; Le Floc'h et al. 2007). With the only exception of CXOJ1417 that is the least luminous sources of the sample, all the remaining cases are characterized by mm luminosities consistent with those of our sources. All the obscured AGNs shown in panels **b**–**d** in Fig. 4 show flat NIR SEDs with similar luminosities. Since the maximum of stellar light in νF_ν and the minimum of AGN light are at NIR wavelengths (1–1.6 μm in the rest-frame; Sawicki 2002; Sanders et al. 1989), it is at these wavelengths that we can expect the maximum contribution to the optical-IR SED from the host galaxy. The similarity of the observed NIR SEDs might be thus explained by a significant contribution from stellar emission. All SEDs are characterized by red MIR colors, consistent with hot dust thermal emission, but such a component is more luminous and redder in our sources than in the sources from the literature. Our sources are also systematically more luminous at optical wavelengths, and this is mainly due to the strong emission lines. In summary, our targets have SEDs that are intermediate between those of type 1 QSOs (panel **a**), and those of obscured QSOs (panels **b**–**d**). In absence of high spatial resolution data and MIR spectroscopy, we cannot determine the origin of the observed differences. Assuming that the intrinsic AGN SED is the same in all sources, the redder MIR SED and the higher optical flux imply a higher AGN luminosity and more extinction

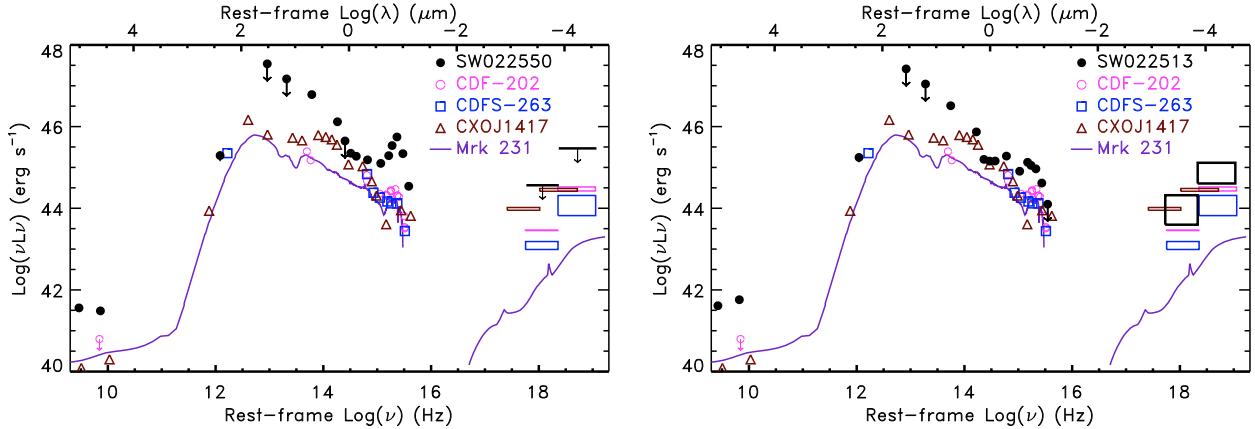


Fig. 5. Rest-frame radio-X ray spectral energy distributions of the selected targets (black full circles), SW022550 (left panel), and SW022513 (right panel) compared with the SEDs of X ray detected heavily absorbed AGNs from the literature, CDF-202 at $z = 3.7$ (magenta circles; Norman et al. 2002; Miller et al. 2008), CDFS-263 at $z = 3.66$ (blue squares; Mainieri et al. 2005), CXOJ1417 at $z = 1.15$ (brown triangles; Le Floc'h et al. 2007), and Mrk 231 at $z = 0.042$ (purple solid line). The rectangles represent the soft and hard X ray fluxes and associated uncertainties. Downward arrows represent 5σ upper limits.

in our sources than in the mm-detected obscured QSOs from the literature.

In Fig. 5, we compare our sources radio-X ray SEDs with those of high- z AGNs with similar multi-wavelength coverage. We include the mm-detected sources CDFS-263, and CXOJ1417 because they are also detected at X ray wavelengths. CXOJ1417 is also detected at radio wavelengths. We also include CDF-202, a X ray detected heavily absorbed type 2 QSO at $z = 3.7$ (Norman et al. 2002), and Mrk 231, a BAL QSO which is heavily absorbed in the X rays and is hosted by a powerful starburst galaxy. We do not consider F 10214+4724 because its X ray spectrum is likely dominated by starburst emission (Alexander et al. 2005b). The absorption-corrected hard X ray luminosities of CDFS-263, CXOJ1417, and CDF-202 are 7.6×10^{44} , 2.4×10^{44} , and 3.3×10^{44} erg s^{-1} , respectively. The X ray luminosity of SW022513 is $\sim 6 \times 10^{44}$ erg s^{-1} . Unless, the X ray luminosity of SW022513 is largely underestimated (see Sect. 2.6), the current estimate is consistent with those measured in the literature sources. In addition to being more luminous at all wavelengths, our sources show a redder MIR emission than the CDFS sources. Interestingly, such a red MIR emission is also observed in other heavily obscured QSOs at high- z which are also characterized by higher MIR/X ray luminosity ratios ($\sim 13\text{--}40$) than classical AGNs (Polletta et al. 2008). For classical AGNs the MIR/X ray luminosity ratios ($\nu L_{6 \mu\text{m}}/L_{2\text{--}10 \text{ keV}}$) range from about 0.3 to 13 (Polletta et al. 2007, 2008), but for SW022513 and CXOJ1417 this ratio is about 60. Such a prominent MIR component might imply that dust heating is more efficient, that the dust covering factor is particularly high, or that the intrinsic X-ray luminosity is underestimated.

3.2. SED modeling and star formation rate

Our targets have optical and MIR properties typical of obscured QSOs, but their large mm fluxes suggest that they are also experiencing intense star forming activity. Indeed, emission from AGN-heated circumnuclear dust is expected to peak around $10\text{--}30 \mu\text{m}$ and not to contribute significantly in the FIR (see e.g. Granato et al. 2004; Hönig et al. 2006; Fritz et al. 2006). A powerful starburst provides a more likely explanation, but such a hypothesis cannot be confirmed because of the lack of detections at rest-frame wavelengths $\lambda = 6\text{--}250 \mu\text{m}$, contrarily to the cases

for F 10214+4724 and CXOJ1417. Although current data do not constrain the origin of the mm emission, AGN model predictions and the similarity with F 10214+4724 and CXOJ1417 favor a starburst origin.

Assuming that the mm flux is due to a starburst component and that the MIR flux is mainly due to the AGN, we model the IR ($2\text{--}1200 \mu\text{m}$) observed SED of our targets by combining a starburst and a torus model. To model the starburst component, we use a library of 7208 starburst models (Siebenmorgen & Krügel 2007). To model the AGN component, we use a set of 44 clumpy torus models with various inclinations from Hönig et al. (2006). These models were used to fit the MIR SEDs and spectra of a sample of MIR-luminous obscured QSOs with available MIR spectra (Polletta et al. 2008). The starburst component is normalized at the observed 1.2 mm flux, and the torus models at the observed $24 \mu\text{m}$ flux. The sum of the starburst and torus models is then compared to the SED of our targets. Since none of the models is able to reproduce the optical data, and since the optical data indicate that there is a contribution from strong emission lines, typical of AGNs, an additional component is added to the starburst+torus models represented by an AGN template. For this component we use a type 1 QSO template (Hatziminaoglou et al. 2005; Polletta et al. 2007) because it reproduces well the observed optical broad-band data. The type 1 QSO template is normalized at the J -band ($1.25 \mu\text{m}$) observed flux. Note that such a normalization corresponds to $\sim 1.4\%$ of what would be obtained if we normalized the type 1 QSO template at the observed $24 \mu\text{m}$ flux. Such a small fraction implies that the AGN optical light might be suppressed by more than 70%, which corresponds to a foreground extinction of $A_V = 4.6$. Since this component is not reddened, its origin is more likely due to a small (1.4%) fraction of nuclear light that escapes without suffering obscuration, or to nuclear scattered light (see e.g. Zakamska et al. 2006).

After combining all the starbursts models, with all the torus models and with the type 1 QSO template, we then reject all combinations that overpredict the fluxes at 70, and $160 \mu\text{m}$. This criterion rejects about 40% of all models. We then select 50 models with the lowest χ^2 , computed by comparing the predicted and observed fluxes in 7 bands, from $1.25 \mu\text{m}$ to $24 \mu\text{m}$. The best 50 models are shown in Fig. 6. The choice of 50 models yields a wide enough range of solutions to characterize all acceptable models and the luminosities produced by each component. A

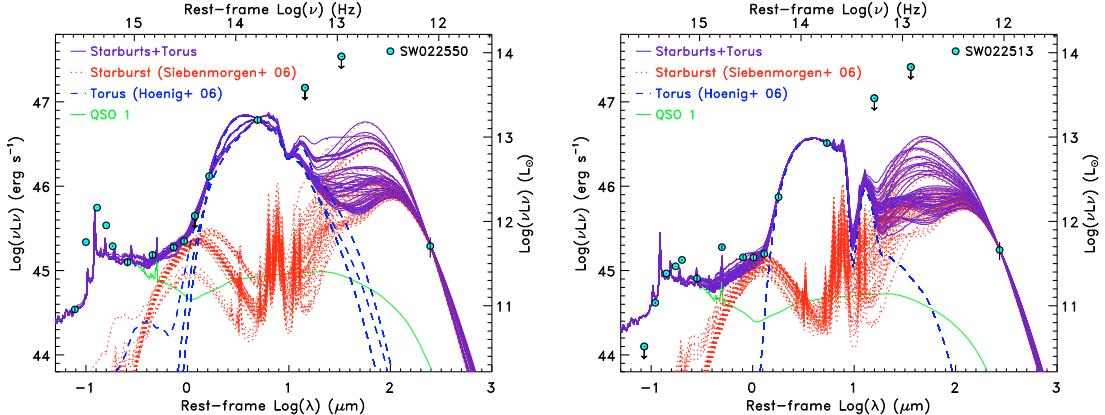


Fig. 6. Rest-frame UV-mm spectral energy distributions of the selected targets (cyan full circles), SW022550 (left panel), and SW022513 (right panel). The red dotted curves represent the starburst models normalized at the observed 1.2 mm flux from the library in Siebenmorgen & Krügel (2007). The dashed blue curves represent torus models from Höning et al. (2006) normalized at 24 μm . The green solid line is a type 1 AGN template normalized at the observed J-band (1.25 μm) flux. The solid purple curves represent the total of the starburst, torus, and type 1 AGN templates which yield the 50 best χ^2 , where the χ^2 is computed using 7 data points from the J-band to the 24 μm , and do not overpredict the observed fluxes and upper limits at FIR wavelengths (see text for more details). The models are used to estimate the luminosities reported in Table 3.

higher number of models yields either poor fits or does not increase significantly the range of SEDs already represented by the 50 best models.

The preferred starburst models contain an old stellar population component, correspond to starbursts of intermediate sizes, i.e. 3 kpc, and show a wide range of ratios, \sim 40–90% of the OB stars luminosity with respect to the total luminosity. The preferred torus model for both sources, defined as the most often chosen among the best 50 models, corresponds to a torus with axis at only 15° from the line of sight with an additional cold dusty absorber along the line of sight. Such a model explains the prominent MIR emission as radiation from hot dust in the torus inner walls, and the red MIR colors as extinguished torus radiation by a cold absorber (for details on the AGN models see Polletta et al. 2008).

From the set of acceptable models we derive the AGN and the starburst bolometric (0.1–1000 μm) luminosities, and the IR (8–1000 μm) starburst luminosity. The derived luminosities for all acceptable models are reported in Table 3.

Assuming that the estimated IR luminosity is produced by a starburst, the derived SFRs, given by $SFR(M_\odot/\text{yr}) = L(\text{IR})/(5.8 \times 10^9 L_\odot)$ (Kennicutt 1998), are \sim 500–3000 M_\odot/yr for both sources. The estimated SFRs are consistent with those estimated in other radio-quiet obscured QSOs at high- z (e.g. Ivison et al. 1998; Efstathiou 2006; Stevens et al. 2005; Polletta et al. 2008; Sajina et al. 2007b).

3.3. The origin of the radio luminosity

Among the various feedback mechanisms that have been suggested as star formation regulators in high- z galaxies, the radio feedback mode is probably the most popular and successful in reproducing numerous observables through simulations (Silk & Rees 1998; Croton et al. 2006; Cattaneo & Teyssier 2007). Because of the predicted short timescale associated with the radio activity and of the difficulty of identifying AGN-driven radio activity in non radio-loud objects, there is only limited evidence of radio feedback at play on star-formation in AGNs, especially at high- z , and at high luminosities. The best supporting observations are the detection of AGN-driven outflows in high- z radio-galaxies (Nesvadba et al. 2006, 2007b,a), and of moderate AGN-driven radio activity in high- z Spitzer-selected

starburst and AGNs (Sajina et al. 2007a; Polletta 2008). Finding further evidence of AGN-driven radio activity, especially in massive high- z star forming galaxies, would provide additional support for the radio feedback as quenching mechanism. Here, we analyze the origin of the radio emission in our two sources.

The measured radio fluxes of SW022550, and SW022513 correspond to 1.4 GHz rest-frame luminosities of $10^{25.50} \text{ W Hz}^{-1}$ and $10^{25.38} \text{ W Hz}^{-1}$, respectively, or $SFRs \sim 6000 M_\odot/\text{yr}$, if powered by star formation. These high luminosities would imply the presence of AGN-driven radio emission in the local universe (an AGN is classified radio-loud if $P_{1.4 \text{ GHz}} > 10^{23.5} \text{ W Hz}^{-1}$, corresponding to a $SFR \simeq 100 M_\odot/\text{yr}$; Condon 1992). However, at high- z , intense starburst episodes with SFRs of thousands of M_\odot/yr can be expected. It is thus not straightforward to infer an AGN origin of the radio emission in sources at high- z based on the estimated radio luminosities. Although a broad multi-wavelength coverage is available for both sources, some traditional diagnostic methods would not be reliable in probing an AGN origin of the radio emission. The commonly diagnostic based on large radio/optical flux ratios cannot be applied because our sources suffer from heavy optical obscuration. Another diagnostic is provided by the radio/infrared flux ratio, which is based on the well established far-IR/radio correlation (Condon 1992). The average FIR-to-radio emission ratio $q = \log(L(\text{FIR})/3.75 \times 10^{12} \text{ W}) - \log(L_{\nu, 1.4 \text{ GHz}}/\text{W Hz}^{-1})$ is equal to 2.34 in local star forming galaxies (Yun et al. 2001), where $L(\text{FIR}) = L_{40-120 \mu\text{m}}$, and $L(\text{FIR})/L(\text{IR}) = 0.56 \pm 0.1$ for the starburst models used in this work. Slightly lower values have been found in sources at higher z , $q = 2.07$ in $z \sim 2$ SMGs (Kovács et al. 2006), and $q = 2.21 \pm 0.02$ in radio-quiet MIR selected $z \sim 1$ – 2.5 AGNs (Sajina et al. 2008). As the AGN contribution to the radio emission increases, q decreases. For example, in radio loud AGNs q is typically close to zero or negative. The q values derived for our sources using the FIR luminosities obtained from the starburst models (see Table 3) are \sim 0.8–1.6 for both sources. These low values favor an AGN origin for the observed radio emission. This is also supported by the steep radio spectral index in SW022550 and the extended radio size of SW022513.

In summary, we find some evidence of an AGN origin for the radio emission of our two sources. The associated radio activity is moderate, as it is only slightly in excess of what is typically

observed in radio quiet quasars and much less than in radio loud quasars.

3.4. AGN and starburst contribution to the bolometric luminosity

The SEDs of our sources indicate that their MIR emission is dominated by AGN-heated dust. Using the best torus models described in Sect. 3.2 and shown in Fig. 6, we estimate a lower limit to the AGN bolometric luminosity by integrating the torus model in the 0.1 – $1000\mu\text{m}$ wavelength range. The derived AGN bolometric luminosities, also reported in Table 3, are $10^{13.4} L_\odot$, and $10^{13.1} L_\odot$ for SW022550 and SW022513, respectively. To derive the starburst bolometric luminosity, we do the same integration but using only the starburst models. Based on these approximations, the AGN contribution to the bolometric luminosities is ~ 50 – 87% in SW022550, and 43 – 83% in SW022513. The broad range of values is due to the uncertainty on the starburst IR luminosity. The AGN contribution to the system bolometric luminosity is thus at least 40%, and it can be more than 80%.

3.5. Host stellar luminosity

The stellar emission, in F_ν , from a host galaxy peaks in the NIR, typically at $1.6\mu\text{m}$ in the rest-frame (or H -band) (Sawicki 2002). The ratio between NIR luminosity and stellar mass is characterized by little dispersion, especially when similar star formation histories are assumed. The estimated stellar masses can be 30% lower in case of younger stellar populations. Thus, we can assume that the NIR luminosity is a proxy of the stellar mass or luminosity of the host galaxy. Since in our objects the AGN optical and NIR emissions are obscured, we can directly estimate the host NIR luminosity from the observed SED. Indeed, the best-fit SED models show that the NIR emission is dominated by stellar light. We estimate the host, and thus stellar, H -band host luminosities using the acceptable starburst templates described above. From the H -band luminosity, we estimate the stellar mass using the average $L(H)/M_*$ ratio derived for a sample of high- z radio galaxies by Seymour et al. (2007), $\text{Log}(M_*) = \text{Log}(L(H)) - 0.1 \pm 0.1$, where M_* is the stellar mass in M_\odot and $L(H)$ is the monochromatic H -band luminosity in L_\odot . Note that this relationship assumes a Kroupa (2001) initial mass function (IMF). The estimated stellar masses are 1.3 – $6.5 \times 10^{11} M_\odot$ for SW022550, and 1.5 – $2.9 \times 10^{11} M_\odot$ for SW022513. The estimated H -band luminosities and stellar masses are reported in Table 3. Because of the uncertainty in separating the host and the AGN contributions to the H -band luminosity, the stellar masses should be considered as upper limits to the true stellar masses. The estimated masses are of the same order as those of the most massive systems at high- z (see e.g. Seymour et al. 2007; Berta et al. 2007), and are thus among the most massive objects at $z \sim 3$ – 4 .

4. Discussion

4.1. Dust obscuration: geometry and distribution

The two sources analyzed in this work exhibit properties usually observed in type 2 AGNs. Their narrow line dominated spectra, low optical/IR flux ratios and faint X ray fluxes imply that they are obscured, but both sources show a blue optical continuum which is unexpected by the standard AGN unification model (Antonucci 1993), but not unusual for obscured AGNs (see e.g. SW104409; Polletta et al. 2006). Moreover,

SW022550 shows some broad components in its optical spectrum. In addition, both sources also exhibit a prominent MIR emission suggesting that the hottest dust component, likely at the dust sublimation radius, is in part visible (see preferred torus models in Sect. 3.2). These hybrid properties suggest that a fraction of the nuclear light is visible. Since the SEDs of our sources cannot be simply reproduced by extinguishing the emission of a type 1 QSO, a more complex model is required to explain their properties. A proper modeling, that takes into account various geometries, dust properties, the intensity and the spectrum and luminosity of the heating source, and a radiative transfer treatment would be necessary to fully explain their hybrid properties. However, the uncertainty on determining the spectrum of the AGN component due mainly to the contribution from the host galaxy, and the lack of a MIR spectrum, would yield degeneracies in the models. Thus, we limit our discussion to a qualitative comparison with some of the most popular models of obscuring dust in AGNs from the literature. A possible scenario is a clumpy dust distribution which produces a partial covering of the nucleus (Nenkova et al. 2002; Hönig et al. 2006). This scenario is quite unlikely if the optical radiation comes from the nuclear region, and it is thus much smaller than the dust sublimation radius. An alternative to an unobstructed line of sight to the nucleus is having a scattering medium, perhaps associated with the NLR. The scattering medium can produce the blue optical light, and broad emission lines. If dust is mixed with the scattering medium, this might be responsible for the large MIR excess, while the nuclear light is obscured by the dusty torus (Schweitzer et al. 2008). Spectropolarimetric observations to measure the scattered component are planned for SW022550 to investigate the latter scenario.

4.2. SMG-QSO evolutionary link

Our targets show simultaneous powerful starburst and AGN activity. They are thus good candidates for transitioning objects between the starburst and the QSO stage predicted by the most popular evolutionary models (Sanders et al. 1988; Alexander et al. 2005c). Here, we consider their properties to make predictions on their evolution.

From the AGN bolometric luminosity (see Table 3), and assuming that the AGN emission is Eddington limited ($L_{\text{bol}} = L_{\text{Edd}}$), we derive a lower limit to the BH masses of SW022550 and SW022513 of $\simeq 7.4 \times 10^8 M_\odot$ and $3.7 \times 10^8 M_\odot$, respectively. Since most of the quasars accrete below their Eddington limit (McLure & Dunlop 2004), it would be more realistic to assume a lower Eddington ratio, but this would imply even higher BH masses. Note that these estimates are not quite accurate since we are estimating the AGN bolometric luminosity from the re-processed thermal emission which is highly dependent on the properties of the reprocessing dust (see e.g. Marconi et al. 2004). Assuming an accretion efficiency of 10%, the derived accretion rates are $\sim 17 M_\odot \text{ yr}^{-1}$ for SW022550 and $8.5 M_\odot \text{ yr}^{-1}$ for SW022513.

Assuming the estimated BH masses and the stellar masses derived in Sect. 3.5, we find that these sources lie on the local $M_{\text{BH}} - M_{\text{bulge}}$ relationship (e.g. Marconi & Hunt 2003). Thus they are more similar to other obscured QSOs at high- z , than to SMGs with AGN activity or type 1 AGNs which are both offset from the local $M_{\text{BH}} - M_{\text{bulge}}$ relationship (Alexander et al. 2008; Polletta 2008; Coppin et al. 2008). In order to keep these objects on the local $M_{\text{BH}} - M_{\text{bulge}}$ relation as they evolve, the on-going star formation process and BH growth will have to stop, or continue with analogous rates ($SFR \equiv \dot{M}$). Since accretion rates

of thousands, or even hundreds of M_{\odot}/yr are not sustainable, the latter scenario can occur only if the SFR decreases significantly. In either case, we predict a quick decrease in the SFR of our sources. The presence of shocks and outflowing gas and the moderate AGN-driven radio activity might be the expression of an AGN-induced feedback that will cause the predicted SFR to decrease. Such a mechanism is predicted in numerous current evolutionary models (Silk 2005; Croton et al. 2006), but it is more often invoked in AGNs at lower redshifts and luminosities. Our results suggest that radio feedback might be also important in high- z QSOs (see also Sajina et al. 2008; Polletta 2008).

5. Summary

We analyze the multi-wavelength spectral energy distributions of two obscured QSOs at high- z discovered in the CFHTLS-D1/SWIRE survey, SWIRE4 J022550.67–042142.37 ($z = 3.867$) and SWIRE4 J022513.92–043420.24 ($z = 3.427$). Both sources benefit from multi-wavelength data available in the field and from mm data from MAMBO and optical and NIR spectroscopy obtained as part of a dedicated follow up program.

Their large mm fluxes, $>4\text{ mJy}$, imply IR luminosities $\sim 10^{12.5-13.3} L_{\odot}$, and SFRs $\sim 500-3000 M_{\odot}/\text{yr}$. The two sources also show powerful AGN activity, mostly seen in the MIR ($L(\text{MIR})_{\text{AGN}} \geq 10^{13} L_{\odot}$) and in the optical emission lines, and probably also at radio wavelengths. Their MIR emission is mainly due to AGN-heated hot dust and corresponds to a luminosity comparable to or higher than the starburst luminosity. The MIR SEDs are redder and more luminous than observed in other mm-detected obscured QSOs (see Sect. 3.1 and Fig. 4).

Their optical and NIR spectra are characterized by faint or absent continuum emission and strong emission lines typical of AGNs. The optical spectrum of SW022550 suggests either a high metallicity, $Z \sim 4 Z_{\odot}$, or the presence of shocks. The comparison between the rest-frame ultraviolet (UV) and optical spectra indicates the presence of outflowing material with a velocity of $\sim 500 \text{ km s}^{-1}$. The red and faint NIR emission compared to the MIR emission indicates that the AGN is extinguished by dust with $A_V > 4$. The faint X ray emission, the large hardness ratio, and the small X ray/[O III] $\lambda 5007$, and X ray/MIR luminosity ratios suggest that their X ray emission might be absorbed by a gas column density $\geq 10^{24} \text{ cm}^{-2}$, and thus that they are Compton-thick QSOs.

The sources show some hybrid properties in between those of type 1 and type 2 AGNs, i.e. a broad component in some UV emission lines, blue UV colors, and a significant emission from hot dust, likely at the sublimation radius. We suggest that these properties are explained by the presence of scattered light at UV wavelengths, and that a hot dust component might be visible, perhaps associated with the NLR and with the scattering medium.

The scattering scenario will be tested with planned spectropolarimetric observations on the VLT of SW022550. With those data we will investigate whether the UV emission of SW022550 is dominated by a scattered component, and measure the scattered fraction.

The composite (starburst and AGN) nature of our sources suggests that they might be experiencing the transition phase between the starburst and the QSO stage predicted by some evolutionary models (e.g. Di Matteo et al. 2005). Since both sources already lie on the local $M_{\text{BH}}-M_{\text{bulge}}$ relation, we predict a quick decrease in their SFR in order to remain on such a relation as they evolve. Both sources are characterized by AGN-driven

moderate radio activity, providing support to evolutionary models invoking radio feedback as star formation quenching mechanism even in powerful QSOs (Silk 2005; Croton 2006).

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References

Alexander, D. M., Bauer, F. E., Chapman, S. C., et al. 2005a, ApJ, 632, 736
 Alexander, D. M., Chartas, G., Bauer, F. E., et al. 2005b, MNRAS, 357, L16
 Alexander, D. M., Smail, I., Bauer, F. E., et al. 2005c, Nature, 434, 738
 Alexander, D. M., Brandt, W. N., Smail, I., et al. 2008, AJ, 135, 1968
 Allen, M. G., Groves, B. A., Dopita, M. A., Sutherland, R. S., & Kewley, L. J. 2008, ApJS, 178, 20
 Antonucci, R. 1993, ARA&A, 31, 473
 Baldwin, J. A., McMahon, R., Hazard, C., & Williams, R. E. 1988, ApJ, 327, 103
 Berta, S., Lonsdale, C. J., Polletta, M., et al. 2007, A&A, 476, 151
 Bondi, M., Ciliegi, P., Zamorani, G., et al. 2003, A&A, 403, 857
 Bondi, M., Ciliegi, P., Venturi, T., et al. 2007, A&A, 463, 519
 Cattaneo, A., & Teyssier, R. 2007, MNRAS, 376, 1547
 Chakrabarti, S., Cox, T. J., Hernquist, L., et al. 2007, ApJ, 658, 840
 Chapman, S. C., Blain, A. W., Smail, I., & Ivison, R. J. 2005, ApJ, 622, 772
 Chiappetti, L., Tajer, M., Trinchieri, G., et al. 2005, A&A, 439, 413
 Ciliegi, P., Zamorani, G., Hasinger, G., et al. 2003, A&A, 398, 901
 Ciliegi, P., Zamorani, G., Bondi, M., et al. 2005, A&A, 441, 879
 Condon, J. J. 1992, ARA&A, 30, 575
 Coppin, K. E. K., Swinbank, A. M., Neri, R., et al. 2008, MNRAS, 389, 45
 Croton, D. J. 2006, MNRAS, 369, 1808
 Croton, D. J., Springel, V., White, S. D. M., et al. 2006, MNRAS, 365, 11
 Di Matteo, T., Springel, V., & Hernquist, L. 2005, Nature, 433, 604
 Dickey, J. M., & Lockman, F. J. 1990, ARA&A, 28, 215
 Dye, S., Warren, S. J., Hambly, N. C., et al. 2006, MNRAS, 372, 1227
 Efstathiou, A. 2006, MNRAS, 371, L70
 Elvis, M., Wilkes, B. J., McDowell, J. C., et al. 1994, ApJS, 95, 1
 Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, ApJS, 154, 10
 Ferrarese, L. 2002, ApJ, 578, 90
 Fritz, J., Franceschini, A., & Hatziminaoglou, E. 2006, MNRAS, 366, 767
 Gallagher, J. S., Hunter, D. A., & Bushouse, H. 1989, AJ, 97, 700
 Gebhardt, K., Bender, R., Bower, G., et al. 2000, ApJ, 539, L13
 Goodrich, R. W., Miller, J. S., Martel, A., et al. 1996, ApJ, 456, L9
 Granato, G. L., Silva, L., Monaco, P., et al. 2001, MNRAS, 324, 757
 Granato, G. L., De Zotti, G., Silva, L., Bressan, A., & Danese, L. 2004, ApJ, 600, 580
 Groves, B. A., Dopita, M. A., & Sutherland, R. S. 2004, ApJS, 153, 75
 Hamann, F., & Ferland, G. 1993, ApJ, 418, 11
 Hao, C.-N., Xia, X.-Y., Shu-DeMao, Deng, Z.-G., & Wu, H. 2008, Chinese J. Astron. Astrophys., 8, 12
 Hartwick, F. D. A. 2004, ApJ, 603, 108
 Hasinger, G., Miyaji, T., & Schmidt, M. 2005, A&A, 441, 417
 Hatziminaoglou, E., Pérez-Fournon, I., Polletta, M., et al. 2005, AJ, 129, 1198
 Hines, D. C., Krause, O., Rieke, G. H., et al. 2006, ApJ, 641, L85
 Höning, S. F., Beckert, T., Ohnaka, K., & Weigelt, G. 2006, A&A, 452, 459
 Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2005a, ApJ, 630, 705
 Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2005b, ApJ, 630, 716
 Hopkins, P. F., Hernquist, L., Cox, T. J., et al. 2006, ApJ, 639, 700
 Ivison, R. J., Smail, I., Le Borgne, J.-F., et al. 1998, MNRAS, 298, 583
 Kennicutt, Jr., R. C. 1998, ARA&A, 36, 189
 Kovács, A., Chapman, S. C., Dowell, C. D., et al. 2006, ApJ, 650, 592

Kreysa, E., Gemuend, H.-P., Gromke, J., et al. 1998, Advanced Technology MMW, Radio, and Terahertz Telescopes, ed. T. G. Phillips, Proc. SPIE, 3357, 319

Kroupa, P. 2001, MNRAS, 322, 231

Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, MNRAS, 379, 1599

Le Floc'h, E., Willmer, C. N. A., Noeske, K., et al. 2007, ApJ, 660, L65

Li, Y., Hernquist, L., Robertson, B., et al. 2007, ApJ, 665, 187

Lonsdale, C. J., Polletta, M., Omont, A., et al. 2008, ApJ, accepted [arXiv: 0810.1293]

Lonsdale, C. J., Smith, H. E., Rowan-Robinson, M., et al. 2003, PASP, 115, 897

Mainieri, V., Rigopoulou, D., Lehmann, I., et al. 2005, MNRAS, 356, 1571

Marconi, A., & Hunt, L. K. 2003, ApJ, 589, L21

Marconi, A., Risaliti, G., Gilli, R., et al. 2004, MNRAS, 351, 169

Marulli, F., Bonoli, S., Branchini, E., Moscardini, L., & Springel, V. 2008, MNRAS, 257

McLure, R. J., & Dunlop, J. S. 2004, MNRAS, 352, 1390

Miller, N. A., Fomalont, E. B., Kellermann, K. I., et al. 2008, ApJS, accepted [arXiv: 0804.2375]

Moorwood, A., Cuby, J.-G., Biereichel, P., et al. 1998, The Messenger, 94, 7

Moy, E., & Rocca-Volmerange, B. 2002, A&A, 383, 46

Mulchaey, J. S., Koratkar, A., Ward, M. J., et al. 1994, ApJ, 436, 586

Nagao, T., Maiolino, R., & Marconi, A. 2006, A&A, 447, 863

Nenкова, M., Ivezić, Ž., & Elitzur, M. 2002, ApJ, 570, L9

Nesvadba, N. P. H., Lehnert, M. D., Eisenhauer, F., et al. 2006, ApJ, 650, 693

Nesvadba, N. P. H., Lehnert, M. D., De Breuck, C., Gilbert, A., & van Breugel, W. 2007a, A&A, 475, 145

Nesvadba, N. P. H., Lehnert, M. D., Genzel, R., et al. 2007b, ApJ, 657, 725

Netzer, H., Mainieri, V., Rosati, P., & Trakhtenbrot, B. 2006, A&A, 453, 525

Norman, C., Hasinger, G., Giacconi, R., et al. 2002, ApJ, 571, 218

Osterbrock, D. E. 1989, Astrophysics of gaseous nebulae and active galactic nuclei, Research supported by the University of California, John Simon Guggenheim Memorial Foundation, University of Minnesota (Mill Valley, CA: University Science Books), 422

Page, M. J., Stevens, J. A., Ivison, R. J., & Carrera, F. J. 2004, ApJ, 611, L85

Polletta, M. 2008, A&A, 480, L41

Polletta, M., Wilkes, B. J., Siana, B., et al. 2006, ApJ, 642, 673

Polletta, M., Tajer, M., Maraschi, L., et al. 2007, ApJ, 663, 81

Polletta, M., Weedman, D., Hönig, S., et al. 2008, ApJ, 675, 960

Pope, A., Chary, R.-R., Alexander, D. M., et al. 2008, ApJ, 675, 1171

Richards, G. T., Nichol, R. C., Gray, A. G., et al. 2004, ApJS, 155, 257

Rieke, G. H., Young, E. T., Engelbracht, C. W., et al. 2004, ApJS, 154, 25

Roettgering, H. J. A., Lacy, M., Miley, G. K., Chambers, K. C., & Saunders, R. 1994, A&AS, 108, 79

Rowan-Robinson, M., Efstathiou, A., Lawrence, A., et al. 1993, MNRAS, 261, 513

Sajina, A., Yan, L., Armus, L., et al. 2007a, ApJ, 664, 713

Sajina, A., Yan, L., Lacy, M., & Huynh, M. 2007b, ApJ, 667, L17

Sajina, A., Yan, L., Lutz, D., et al. 2008, ApJ, 683, 659

Sanders, D. B., Soifer, B. T., Elias, J. H., Neugebauer, G., & Matthews, K. 1988, ApJ, 328, L35

Sanders, D. B., Phinney, E. S., Neugebauer, G., Soifer, B. T., & Matthews, K. 1989, ApJ, 347, 29

Sawicki, M. 2002, AJ, 124, 3050

Schweitzer, M., Groves, B., Netzer, H., et al. 2008, ApJ, 679, 101

Seymour, N., Stern, D., De Breuck, C., et al. 2007, ApJS, 171, 353

Sharp, R., Saunders, W., Smith, G., et al. 2006, in Ground-based and Airborne Instrumentation for Astronomy, ed. I. S. McLean, & M. Iye, Proc. SPIE, 6269, 62690G

Shields, J. C., Walcher, C. J., Böker, T., et al. 2008, ApJ, 682, 104

Siana, B., Polletta, M. d. C., Smith, H. E., et al. 2008, ApJ, 675, 49

Siebenmorgen, R., & Krügel, E. 2007, A&A, 461, 445

Silk, J. 2005, MNRAS, 364, 1337

Silk, J., & Rees, M. J. 1998, A&A, 331, L1

Spergel, D. N., Verde, L., Peiris, H. V., et al. 2003, ApJS, 148, 175

Springel, V., White, S. D. M., Jenkins, A., et al. 2005, Nature, 435, 629

Steidel, C. C., Adelberger, K. L., Giavalisco, M., Dickinson, M., & Pettini, M. 1999, ApJ, 519, 1

Stern, D., Moran, E. C., Coil, A. L., et al. 2002, ApJ, 568, 71

Stevens, J. A., Page, M. J., Ivison, R. J., et al. 2005, MNRAS, 360, 610

Surace, J. A., Shupe, D. L., Fang, F., et al. 2005, Spitzer Science Center, California Institute of Technology

Szokoly, G. P., Bergeron, J., Hasinger, G., et al. 2004, ApJS, 155, 271

Tajer, M., Polletta, M., Chiappetti, L., et al. 2007, A&A, 467, 73

Teplitz, H. I., Armus, L., Soifer, B. T., et al. 2006, ApJ, 638, L1

Vanden Berk, D. E., Richards, G. T., Bauer, A., et al. 2001, AJ, 122, 549

Vernet, J., Fosbury, R. A. E., Villar-Martín, M., et al. 2001, A&A, 366, 7

Wang, R., Carilli, C. L., Beelen, A., et al. 2007, AJ, 134, 617

Weedman, D., Polletta, M., Lonsdale, C. J., et al. 2006, ApJ, 653, 101

Yan, L., Sajina, A., Fadda, D., et al. 2007, ApJ, 658, 778

Yun, M. S., Reddy, N. A., & Condon, J. J. 2001, ApJ, 554, 803

Zakamska, N. L., Strauss, M. A., Krolik, J. H., et al. 2006, AJ, 132, 1496