

VLA H53 α RADIO RECOMBINATION LINE OBSERVATIONS OF THE ULTRALUMINOUS INFRARED GALAXY ARP 220

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ABSTRACT

We present high angular resolution (0."7) observations made with the Very Large Array (VLA) of the radio recombination line (RRL) H53 α and radio continuum emission at 43 GHz from the ultraluminous infrared galaxy (ULIRG) Arp 220. The 43 GHz continuum emission shows a compact structure ($\sim 2''$) with two peaks separated by $\sim 1''$, the east (E) and west (W) components, that correspond to each galactic nucleus of the merger. The spectral indices for both the E and W components, using radio continuum images at 8.3 and 43 GHz are typical of synchrotron emission ($\alpha \sim -1.0$). Our 43 GHz continuum and H53 α line observations confirm the flux densities predicted by the models proposed by Anantharamaiah et al. (2000). This agreement with the models implies the presence of high-density ($\sim 10^5 \text{ cm}^{-3}$) compact H II regions ($\sim 0.1 \text{ pc}$) in Arp 220. The integrated H53 α line emission is stronger toward the nonthermal radio continuum peaks, which are also coincident with the peaks of molecular emission of the H₂CO. The coincidence between the integrated H53 α and the H₂CO maser line emission suggests that the recent star-forming regions, traced by the high-density gas, are located mainly in regions that are close to the two radio continuum peaks. A velocity gradient of $\sim 0.30 \text{ km s}^{-1} \text{ pc}^{-1}$ in the H53 α RRL is observed toward the E component, and a second velocity gradient of $\sim 0.15 \text{ km s}^{-1} \text{ pc}^{-1}$ is detected toward the W component. The orientations of these velocity gradients are in agreement with previous CO, H I, and OH observations. The kinematics of the high-density ionized gas traced by the H53 α line are consistent with two counterrotating disks as suggested by the CO and H I observations.

Subject headings: galaxies: general — galaxies: individual (Arp 220) — galaxies: starburst — radio continuum: galaxies — radio lines: galaxies

1. INTRODUCTION

Arp 220 (IC 4553/4, UGC 9913, IRAS 15327+2340) has an infrared luminosity of $\sim 1.5 \times 10^{12} L_\odot$, as determined from the *Infrared Astronomical Satellite (IRAS)* 25–100 μm observations (Soifer et al. 1984) and therefore is classified as an ULIRG. At a distance of $\sim 70 \text{ Mpc}$ ($H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$), Arp 220 is characterized in the optical by tidal tails interpreted as the signature of a merging process (Surace et al. 2000). The large visual extinction toward Arp 220, which could be $A_V > 100 \text{ mag}$ (González-Alfonso et al. 2004), precludes the observations of the nuclear regions at optical and even at infrared frequencies. Radio frequency observations do not suffer from dust extinction, thus providing a powerful tool to investigate the kinematics, morphology, and physical properties of the gas in the dust-obscured regions of Arp 220. Radio observations reveal a double-nucleus system separated by $\sim 1''$ ($\sim 350 \text{ pc}$) as observed in CO and H I (Mundell et al. 2001; Sakamoto et al. 1999). The existence of multiple luminous radio supernovae (RSNe) in the nuclear regions of Arp 220 suggests that it is mainly starburst-powered (Smith et al. 1998). However, compact OH maser and X-ray emission suggests the presence of active galactic nucleus (AGN)

activity in the W nucleus (Norris et al. 1985; Lonsdale et al. 1998; Clements & Baker 1996; Clements et al. 2002).

High angular resolution (0."5) CO observations carried out by Sakamoto et al. (1999) suggest that the two nuclei are counterrotating with respect to each other and are embedded in a kiloparsec-size gas disk, which rotates around their dynamical center. The counterrotation is also observed in higher angular resolution (0'22) observations of H I absorption by Mundell et al. (2001), who showed that the rotation is not coplanar and proposed the model in which the two nuclei are in the final state of merging. Counterrotation may provide the mechanism to get rid of angular momentum, a condition for the two nuclei to merge. The senses of the velocity gradients, as determined from OH maser emission observations (Rovilos et al. 2003), are consistent with CO molecular emission and H I absorption observations (Sakamoto et al. 1999; Mundell et al. 2001).

Arp 220 is the most distant system from which radio recombination lines (RRLs) have been observed (Zhao et al. 1996; Anantharamaiah et al. 2000). VLA observations of the centimeter wavelength RRL H92 α have been reported by Zhao et al. (1996) and Anantharamaiah et al. (2000) with angular resolutions of 4" and 1", respectively. Anantharamaiah et al. (2000), using a single-dish telescope (IRAM 30 m), detected the emission of the millimeter-wavelength RRLs H42 α , H40 α , and H31 α , and with the VLA established upper limits for the RRLs H167 α and H165 α . Anantharamaiah et al. (2000) constructed a model using three density components of ionized gas in order to reproduce the observed RRL and radio continuum intensities, obtaining a star formation rate (SFR) in Arp 220 of $\sim 240 M_\odot \text{ yr}^{-1}$. The existence of a high-density component of ionized gas, as deduced from these models, has been interpreted as evidence for recent star formation in Arp 220. In particular,

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TABLE 1
OBSERVATIONAL PARAMETERS FOR ARP 220

Parameter	H53 α Line
Right ascension (J2000.0)	15 ^h 34 ^m 57 ^s .28
Declination (J2000.0)	23°30'11".9
Angular resolution	0".7
Total observing duration (hr)	13
Bandwidth (MHz)	150
Number of spectral channels	45
Center V_{Helio} (km s $^{-1}$)	5500
Velocity coverage (km s $^{-1}$)	1000
Velocity resolution (km s $^{-1}$)	44
Amplitude calibrator	J1331+305
Phase calibrator	J1540+147
Bandpass calibrator	J1229+020
rms line noise per channel (mJy beam $^{-1}$)	1
rms continuum (mJy beam $^{-1}$)	0.3

the RRL H53 α can be used as a tracer of thermal high-density gas ($\sim 10^5$ cm $^{-3}$). The star formation in this system may be supported by multiple starbursts with a high SFR.

In this paper we present VLA observations of Arp 220 in the millimeter regime of the RRL H53 α (7 mm), as well as radio continuum at 43 GHz with subarcsecond (0".7) angular resolution. The observations are described in § 2. In § 3 the results obtained from the 43 GHz radio continuum and the RRL H53 α are presented. In § 4 the implications of these results are discussed. The conclusions are presented in § 5.

2. VLA OBSERVATIONS OF THE H53 α LINE

The RRL H53 α (43 GHz) was observed with the VLA in the C configuration on 2002 October 13 and 25. The flux density scales were determined using observations of J1331+305 (1.49 Jy). The fast-switching mode was used in order to track the phase variations induced by the troposphere using cycles of 230 s on Arp 220 and 40 s on the phase reference source J1540+147 (~ 1 Jy). The bandpass response of the instrument is frequency-dependent, and observations of a bandpass calibrator were required for each frequency window. This bandpass response was corrected using observations of J1229+020 (17.5 ± 0.5 Jy). Five frequency windows, each with 15 spectral channels (~ 22 km s $^{-1}$ each), were used to observe the broad RRL. The total velocity coverage is ~ 1000 km s $^{-1}$. The different frequency windows were centered at 42114.9, 42135.1, 42164.9, 42185.1, and 42214.9 MHz. The on-source integration time was ~ 2 hr for each frequency window. To avoid possible phase decorrelation at 7 mm, the data were calibrated correcting for the phases in a first step and subsequently correcting for both amplitude and phase. The spectral-line data were further calibrated by applying the solutions obtained from the self-calibration performed on the continuum channel (which contains the average of the central 75% of the bandpass) of each frequency window. The spectral-line observations of the second day were not combined with those of the first day since the variable content of water vapor in the troposphere caused decorrelation of phases in the second and third frequency windows. A critical step in the reduction of the data is to determine the bandpass shape by normalizing the bandpass using the channel 0 data. Five line images, one for each frequency window of the first day, were produced using a weighting scheme intermediate between natural and uniform to obtain a circular beam of 0".7. The five line cubes were regridded in frequency and combined into a single line cube in the GIPSY environment. The method used to combine different windows is described by Rodriguez-Rico

et al. (2004). The data were then Hanning-smoothed to minimize the Gibbs effect, and a final velocity resolution of ~ 44 km s $^{-1}$ was achieved. The parameters of the observations are given in Table 1. The resulting line cube was further processed using `imlin` from AIPS to subtract the continuum emission. A first-order polynomial fit was used for the line-free channels, two on the low-velocity edge and five on the high-velocity edge of the bandpass. Previous Arp 220 observations of the RRL H42 α (Anantharamaiah et al. 2000), CO (2–1) emission line (Sakamoto et al. 1999), and H α absorption (Mundell et al. 2001) show that the full width at zero intensity expected for the H53 α line ranges from ~ 5000 to 5700 km s $^{-1}$. The velocity range covered by the H53 α line observations is 5100–6100 km s $^{-1}$, consequently loosing the continuum that corresponds to the blueshifted side of the line. We performed a test using the line-free channels on the blueshifted side of the line as a baseline. We have determined that if weak H53 α line emission with intensity < 3 mJy (3σ) is present in the velocity range of 5000–5100 km s $^{-1}$, the FWHM is not affected. The fit for the H53 α line is in agreement with the fit obtained using the continuum for the low-velocity edge of the bandpass. The H53 α (43 GHz) data have been compared with the RRL H92 α (8.3 GHz) observations taken by Anantharamaiah et al. (2000) with the VLA in the B configuration.

3. RESULTS

3.1. Radio Continuum

Figure 1 shows the continuum image of Arp 220 at 43 GHz superposed on the continuum image at 8.3 GHz, both images with an angular resolution of 0".7. In this image, the two continuum peaks (E and W) are clearly resolved at both frequencies. Using the task `imfit` in AIPS, we estimated the properties of these two components. The physical parameters at 8.3 and 43 GHz were obtained by integrating over equivalent areas in order to estimate the spectral index α ($S_\nu \propto \nu^\alpha$) of each component.

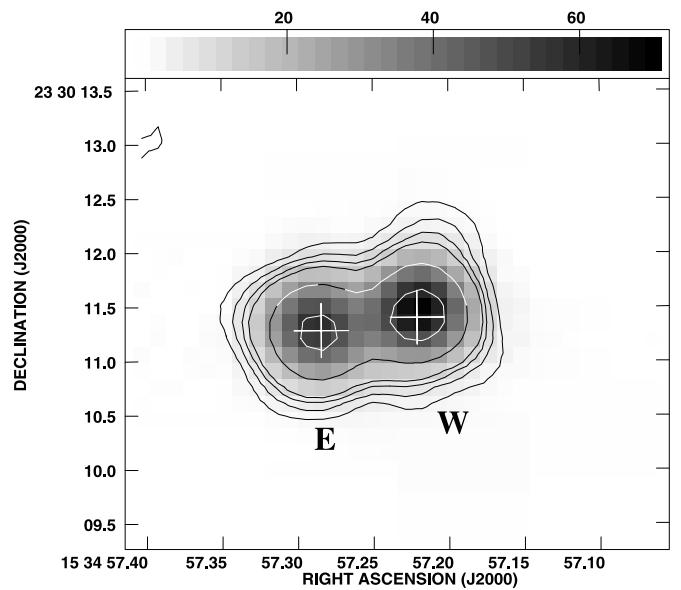


FIG. 1.—Radio continuum image of Arp 220 at 43 GHz (contours) made using VLA observations in the C configuration superposed on the continuum image at 8.3 GHz (gray scale) using the VLA in the B configuration, both images made with angular resolution of 0".7. The contour levels (43 GHz) are drawn at $-3, 3, 5, 7, 9, 18$, and 36 times the rms noise 0.3 mJy beam $^{-1}$. The gray scale (8.3 GHz) ranges from 0.15 to 70 mJy beam $^{-1}$. The crosses mark the peak positions (given in Table 2) of the 43 GHz continuum. The size of the crosses is $0".5$, which is about 5 times the error position of the radio continuum peaks.

TABLE 2
RESULTS FROM GAUSSIAN FITTING TO THE CONTINUUM EMISSION AT 43 GHz OF ARP 220

Feature	R.A. (J2000.0)	Decl. (J2000.0)	Size ^a (arcsec)	P.A. (deg)	S_C (mJy)
Arp 220 E	15 34 57.28 \pm 0.01	23 30 11.3 \pm 0.1	0.5 \times 0.3,	89	17 \pm 2
Arp 220 W	15 34 57.22 \pm 0.01	23 30 11.4 \pm 0.1	0.4 \times 0.2,	114	21 \pm 2

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds.

^a Deconvolved angular size.

TABLE 3
RESULTS FROM GAUSSIAN FITTING TO THE CONTINUUM EMISSION AT 8.3 GHz OF ARP 220

Feature	R.A. (J2000.0)	Decl. (J2000.0)	Size ^a (arcsec)	P.A. (deg)	S_C (mJy) ^b
Arp 220 E	15 34 57.29 \pm 0.01	23 30 11.3 \pm 0.1	0.5 \times 0.4	88	77 \pm 1
Arp 220 W	15 4 57.22 \pm 0.01	23 30 11.5 \pm 0.1	0.5 \times 0.3	107	89 \pm 1

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. Observations at 8.3 GHz made by Anantharamaiah et al. (2000).

^a Deconvolved angular size.

^b Continuum flux densities were measured in the same area as the 43 GHz continuum flux densities listed in Table 2.

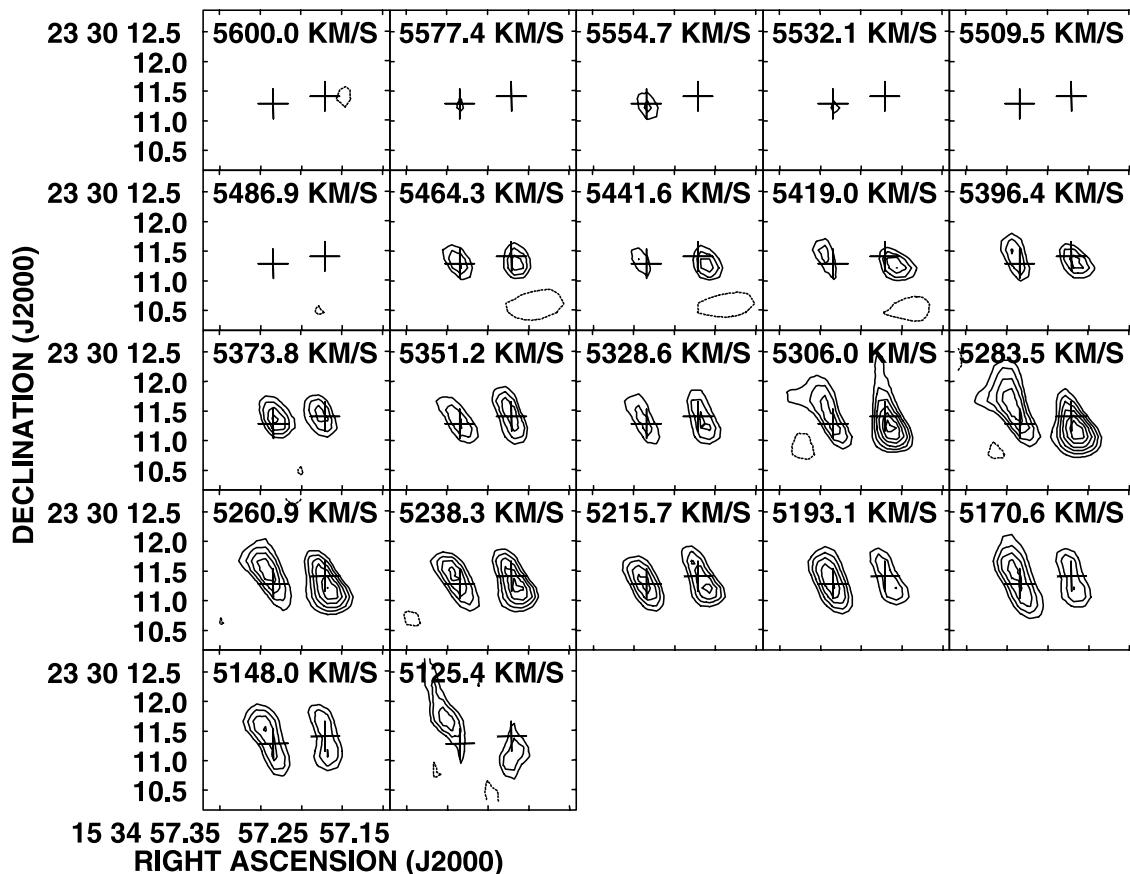


FIG. 2.—H53 α continuum-subtracted channel images from Arp 220 labeled according to the central heliocentric velocity of each channel. Contours are $-3, 3, 4, 5, 6, 7, 8, 9, 10$, and 11 times 1 mJy beam^{-1} , the rms noise. The crosses show the position of the two 43 GHz continuum peaks. The synthesized beam is $0''.7$ FWHM.

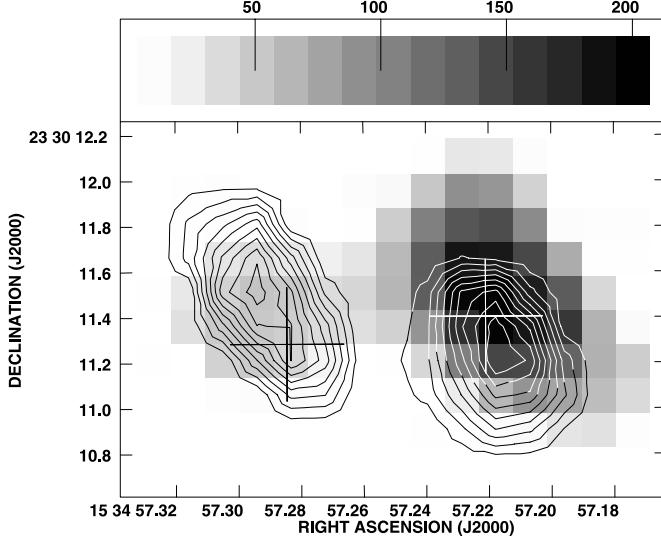


FIG. 3.—Integrated H53 α line emission (moment 0) from Arp 220 in contours with integrated H92 α line emission (moment 0) superposed in gray scale. Contour levels are 10%, 30%, 50%, 70%, and 90% of the peak (1.3 Jy beam $^{-1}$ km s $^{-1}$). The gray scale covers the range 0–200 Jy beam $^{-1}$ km s $^{-1}$. The crosses show the position of the two 43 GHz continuum peaks. The synthesized beam for both RRLs is 0 $''$ 7 FWHM.

Table 2 lists the physical parameters of the continuum components E and W determined at 43 GHz. The total measured 43 GHz continuum flux density of Arp 220 is 44 ± 4 mJy. The E and W components have continuum flux densities of ~ 17 and ~ 21 mJy, respectively. There is also a contribution of ~ 6 mJy to the total flux density from extended emission that was determined using the task `iring` in the AIPS environment by integrating the flux density over an angular scale of 3 $''$ –8 $''$. This total continuum flux density at 43 GHz is in agreement with the value obtained by interpolating from previous observations at 32 GHz and 97 GHz (Downes & Solomon 1998; Anantharamaiah et al. 2000). The continuum physical parameters at 8.3 GHz of both the E and W components in Arp 220 are listed in Table 3. The positions of each continuum peak are coincident at 8.3 and 43 GHz. The two radio nuclei of Arp 220 are separated by $0\farcs9 \pm 0\farcs1$ at a position angle (P.A.) of 95° (east with respect to north). The spectral index between these two frequencies is -0.96 ± 0.08 for the E component and -0.97 ± 0.07 for the W component. These spectral indices are typical of nonthermal emission.

3.2. Radio Recombination Line H53 α

We detected H53 α RRL emission (>3 mJy) toward both components E and W of Arp 220 in the velocity range 5130–5600 km s $^{-1}$. The velocity range over which the H53 α line emission is observed agrees with observations of the RRLs H40 α , H42 α , and H92 α (Anantharamaiah et al. 2000) and the CO (2–1) line (Sakamoto et al. 1999). Figure 2 shows the VLA spectral channel images of H53 α line emission observed over the central ~ 500 pc of Arp 220. Figure 3 shows the velocity-integrated line emission (moment 0) of the H53 α line superposed on moment 0 of the H92 α line. The peaks in moment 0 of the H53 α line emission are nearly coincident ($<0\farcs5$) with the two radio continuum nuclei. Moment 0 of the H92 α line emission (gray scale) exhibits two peaks that are close to the continuum peaks and are coincident with the H53 α peaks within 0 $''$.2.

The total integrated spectrum over the region with detectable H53 α line emission is shown in Figure 4, along with the E and

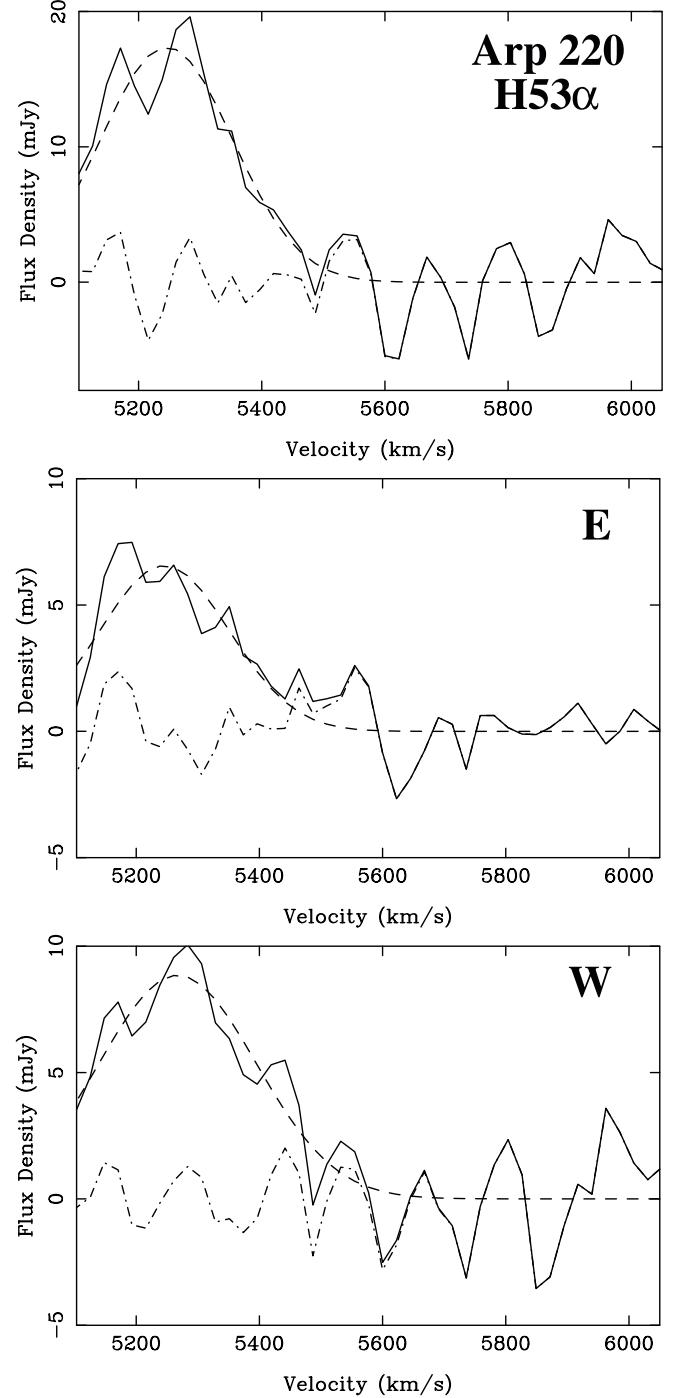


FIG. 4.—H53 α recombination line profiles in Arp 220 observed using the VLA integrated over the region shown in contours in Fig. 1 (top), over the E component (middle), and over the W component (bottom). The dashed line shows the resulting Gaussian fit to the data, and the dash-dotted line indicates the residuals from the Gaussian fit.

W spectra obtained by integrating over each of these components. A single Gaussian fit was carried out for each line profile. The results from the H53 α line fits are listed in Table 4. The peak line flux density is given in column (2), the FWHM of the line in column (3), the heliocentric velocity of the line center in column (4), and the integrated line emission in column (5). Based on the fit parameters for the E and W components, the peak H53 α line flux density for the W component is approximately 30% stronger than the E component. The H53 α line

TABLE 4
RESULTS FROM GAUSSIAN FITTING TO THE H53 α LINE EMISSION OF ARP 220

Feature (1)	S_P (mJy) (2)	ΔV_{FWHM} (km s $^{-1}$) (3)	V_{Helio} (km s $^{-1}$) (4)	$1.07 (S_P \Delta V_{\text{FWHM}})$ (10^{-23} W m $^{-2}$) (5)
Arp 220 E	6 ± 1	235 ± 20	5240 ± 10	220 ± 40
Arp 220 W	8 ± 1	265 ± 30	5270 ± 10	320 ± 50
Arp 220 E+W	16 ± 2	230 ± 20	5250 ± 10	560 ± 80

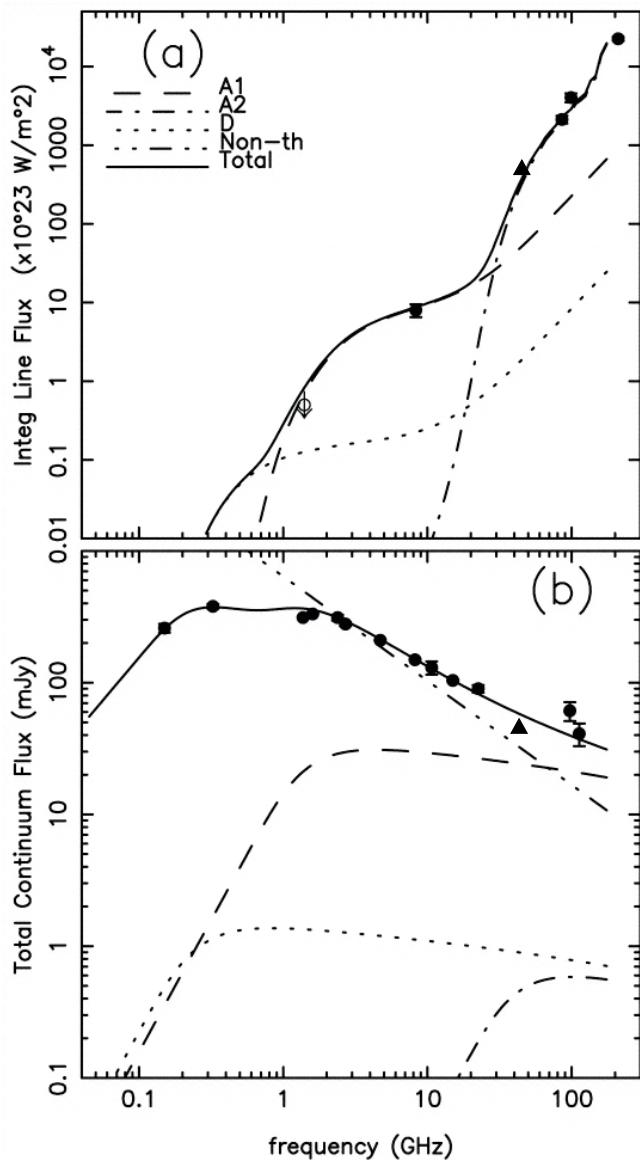


FIG. 5.—Continuum and RRL model results for three ionized gas components (A1, A2, and D) in Arp 220 by Anantharamaiah et al. (2000). The recombination line and continuum flux densities are shown in (a) and (b), respectively. The contribution from the different components are shown and the parameters for each component are given in Table 9 of Anantharamaiah et al. (2000). The observed line data points are listed in Table 2 and 4 of Anantharamaiah et al. (2000). The filled triangles indicate the (a) velocity-integrated H53 α line flux density and (b) 43 GHz continuum flux density obtained in the current observations. The size of the triangles represent the 1σ error for the measured values.

widths and central velocities are similar for the E and W components. The H53 α and H92 α line widths and central velocities determined for the total integrated line emission over Arp 220 are in good agreement with each other. Figure 5 shows the results from radio continuum and RRL emission models taken from Anantharamaiah et al. (2000). The measurements for the total integrated 43 GHz continuum and H53 α line emission are shown in Figures 5a and 5b, respectively.

Figure 6 shows the H53 α line velocity field (moment 1) for the ionized gas in Arp 220. Based on this H53 α velocity field, two velocity gradients are observed toward the E and W components. The velocity gradient on the E component is $\sim 0.30 \pm 0.10$ km s $^{-1}$ pc $^{-1}$, oriented along a P.A. of $\sim 30^\circ$. The velocity gradient on the W component has been marginally detected ($\sim 0.15 \pm 0.10$ km s $^{-1}$ pc $^{-1}$) along a P.A. of $\sim 260^\circ$. Compared to the velocity resolution of the H53 α line (44 km s $^{-1}$), a detailed comparison of the H53 α and the H92 α velocity fields is not possible due to the lower velocity resolution of the H92 α line (230 km s $^{-1}$).

4. DISCUSSION

The measured separation between the two continuum peaks ($0.9 \pm 0.1 \simeq 320$ pc) and the P.A. of their separation (95°) at 43 GHz are in good agreement with previous measurements (~ 350 pc) at 5, 15, 23 (Norris 1988), and 4.8 GHz (Baan & Haschick 1995). The peak positions at 8.3 and 43 GHz agree at the level of 0.1 . The spectral index for Arp 220, determined using the radio continuum flux density of ~ 170 mJy measured at 8.3 GHz and the measurement at 43 GHz of ~ 44 mJy is $\alpha \sim -1.0$. On the other hand, the spectral index determined using the radio continuum flux density of ~ 49 mJy measured at 32 GHz (Baan et al. 1986) and the measurement at 43 GHz (this work) of ~ 44 mJy, is $\alpha \simeq 0.36 (\pm 0.05)$. Anantharamaiah et al. (2000) estimated that the thermal radio continuum flux density at 5 GHz is ~ 32 mJy. If we use a spectral index value $\alpha = -0.1$ for the thermal free-free emission, the expected thermal flux density at 43 GHz is ~ 25 mJy, which is $\sim 60\%$ of the total measured continuum flux density at 43 GHz (see Fig. 5). The flattening of the spectral index, a value of α closer to zero when the 32 and 43 GHz measurements are used, indicates that the ratio between the thermal emission and nonthermal emission is larger at 43 GHz than at 8.3 GHz.

Anantharamaiah et al. (2000) modeled the emission of radio continuum and RRL observations at different frequencies. In these models Anantharamaiah et al. (2000) used three thermally ionized gas components (A1, B1, and D) and nonthermal emission. Components A1 and D are characterized by electron densities of 10^3 cm $^{-3}$, while component B1 has a larger electron density of 2.5×10^5 cm $^{-3}$. All parameters for each of these components are listed in Table 9 of Anantharamaiah et al. (2000). Figure 5 shows the model results from Anantharamaiah et al. (2000), including the measurements of the H53 α line ($\sim 510 \times 10^{-23}$ W m $^{-2}$ or 3.7 Jy km s $^{-1}$) and the continuum flux density at

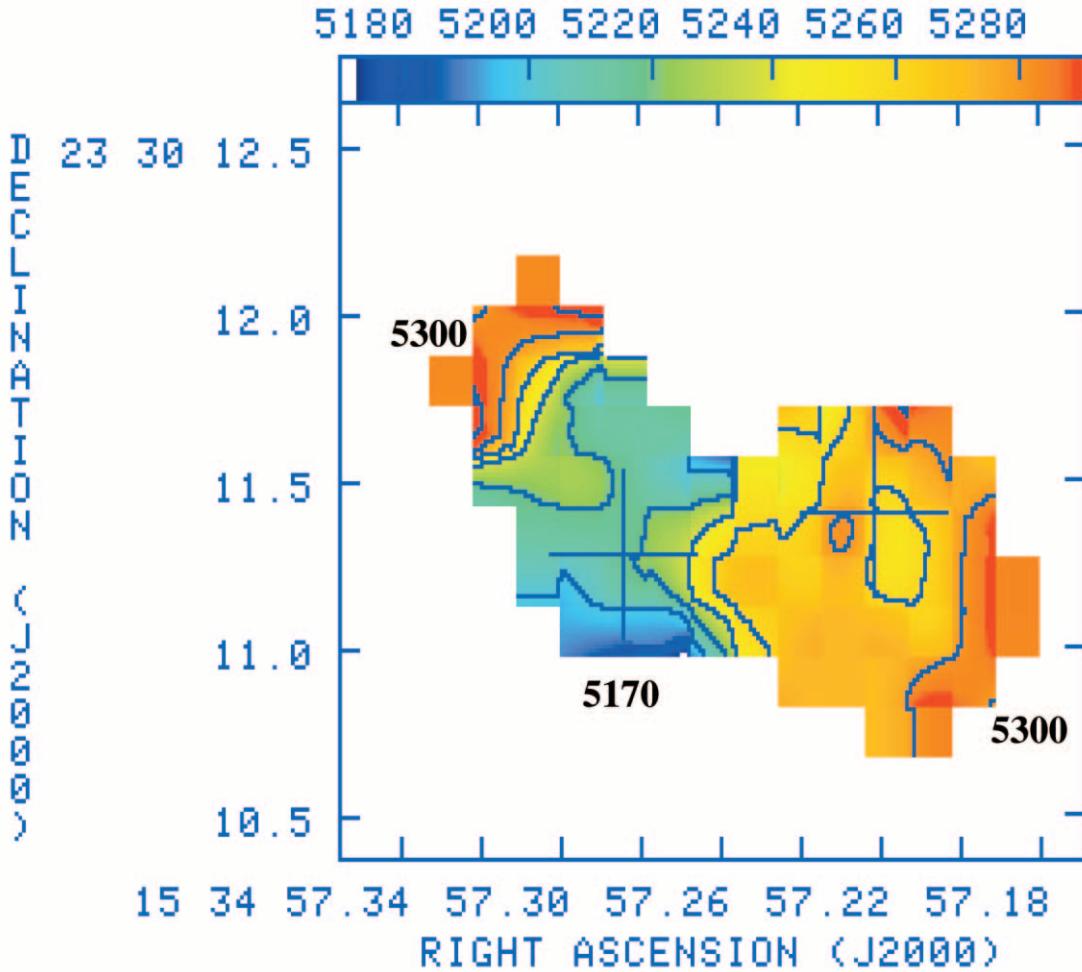


FIG. 6.—Velocity field of Arp 220 as observed in the RRL H53 α with the VLA. The color scale shows the velocity field in the range from 5170 to 5370 km s $^{-1}$. Contour levels are 5180, 5200, . . . , 5360 km s $^{-1}$ in steps of 20 km s $^{-1}$. The crosses show the position of the two 43 GHz continuum peaks. The synthesized beam is 0''.7 FWHM.

43 GHz (\sim 44 mJy). The H53 α line and 43 GHz continuum flux densities are in agreement with the values predicted in the Anantharamaiah et al. (2000) models, confirming the presence of a high-density ionized gas component (\sim 10 5 cm $^{-3}$) composed of compact H II regions with \sim 0.1 pc diameter. According to these models, the contribution from the dense gas to the H92 α line emission is negligible, and the contribution from the lower density gas to the H53 α line is $<10\%$.

On the other hand, the models of Anantharamaiah et al. (2000) indicate that \sim 50% of the H53 α line emission in the high-density ionized gas is due to internal stimulated line emission. The spatial distribution of the H53 α and the H92 α lines is similar toward the E component, as can be seen in Figure 3. However, a remarkable difference is observed toward the W component, where the H92 α line emission is more extended than the H53 α line emission. Near the continuum peaks, the spatial distribution of the H53 α and H92 α RRLs emission suggests that the recent star formation is taking place mostly in two regions. These two regions of line emission are slightly displaced from the respective continuum peaks: the W line peak is $0''.2 \pm 0''.1$ to the south of its associated continuum peak, and the E line peak is $0''.3 \pm 0''.1$ to the north of its associated continuum peak. According to the integrated H53 α line emission (see Fig. 3), the bulk of high-density ionized gas is concentrated in these two regions. This result is supported by the spatial correlation between the two velocity-integrated H53 α line

emission peaks and the peaks of the formaldehyde (H₂CO) molecular emission (Baan & Haschick 1995).

The models of Anantharamaiah et al. (2000) derived a SFR in Arp 220 of $240 M_{\odot}$ yr $^{-1}$, assuming a mass range of 1 – $100 M_{\odot}$ in the Miller-Scalo initial mass function (IMF). According to these models, the SFR could be as low as $90 M_{\odot}$ yr $^{-1}$ if the upper mass limit in the IMF is reduced to $60 M_{\odot}$ and the Salpeter IMF is used. Given that the radio continuum and the velocity-integrated line flux density are in agreement with the values expected from the models of Anantharamaiah et al. (2000), the value for the SFR could be in the range of 90 – $240 M_{\odot}$ yr $^{-1}$, and the mass of high-density ionized gas (\sim 10 5 cm $^{-3}$) is between 10^3 and $10^4 M_{\odot}$.

The orientations of the H53 α velocity gradients in Arp 220, with P.A. \sim 30° in the E component and P.A. \sim 260° in the W component agree with the CO and H I observations (Sakamoto et al. 1999; Mundell et al. 2001). Based on the P.A. of these two velocity gradients, the CO (2–1) line observations were interpreted by Sakamoto et al. (1999) as evidence of two counter-rotating disks in Arp 220. An alternative model that consists of a warped gas disk that resulted from a merger of two spiral galaxies was also proposed by Eckart & Downes (2001). The angular resolution achieved in the H53 α line observations (0''.7) is insufficient to discern between these two models. The velocity structure of the ionized gas based on the H53 α data shows two

velocity gradients in the two components of Arp 220. On the E component the $\text{H}53\alpha$ velocity gradient is $\sim 0.30 \text{ km s}^{-1} \text{ pc}^{-1}$. The H I velocity gradient on the E component is $1.01 \pm 0.02 \text{ km s}^{-1} \text{ pc}^{-1}$ (Mundell et al. 2001), and the OH velocity gradient is $0.32 \pm 0.03 \text{ km s}^{-1} \text{ pc}^{-1}$ (Rovilos et al. 2003).

The velocity gradient on the E component implies a virial mass of $\sim 8 \times 10^7 (\sin^{-2} i) M_\odot$ in a disk of radius $\sim 180 \text{ pc}$ (i is the inclination). The total mass calculated from OH observations, for a disk of radius $\sim 80 \text{ pc}$, is $\sim 10^7 M_\odot$ (Rovilos et al. 2003). The value derived from the OH observations is consistent with the $\text{H}53\alpha$ line observations when the different sizes are taken into account. Based on H I and CO (Sakamoto et al. 1999; Mundell et al. 2001) observations, the estimates of the total mass are $\sim 10^9 M_\odot$. The total mass estimated using CO and H I observations is 2 orders of magnitude larger than the total mass calculated from these $\text{H}53\alpha$ line observations. The smaller estimates obtained from the $\text{H}53\alpha$ line and OH observations are explained if the molecular CO and neutral H I gas are distributed over a larger region compared to the ionized gas. In the direction of the W component, the $\text{H}53\alpha$ velocity gradient is a factor of 10 less than that observed in H I absorption and OH maser emission (Mundell et al. 2001; Rovilos et al. 2003). However, the orientation and sense of the velocity gradient in $\text{H}53\alpha$ are consistent with those determined from the CO, H I , and OH observations (Sakamoto et al. 1999; Mundell et al. 2001; Rovilos et al. 2003).

5. CONCLUSIONS.

We have observed the RRL $\text{H}53\alpha$ and radio continuum at 43 GHz toward the ULIRG Arp 220 with high angular resolution ($0.^{\prime\prime}7$) using the VLA. The 43 GHz radio continuum and the RRL $\text{H}53\alpha$ have been compared with observations of the 8.3 GHz radio continuum and $\text{H}92\alpha$ line (Anantharamaiah et al. 2000).

1. The total 43 GHz radio continuum flux density of Arp 220 is $44 \pm 4 \text{ mJy}$. The morphological characteristics observed in

the radio continuum at 43 GHz agree with previous radio observations made in the range from 1.4 to 23 GHz. In the radio continuum at 43 GHz, Arp 220 exhibits a double-nucleus system that has been clearly resolved with a separation of the two radio continuum peaks of $\sim 1''$ at P.A. of 95° .

2. We have determined the spectral indices for both the E and W components, using radio continuum flux densities at 8.3 and 43 GHz. Both components have spectral indices typical of synchrotron emission ($\alpha \sim -0.9$), as expected from extrapolation from the centimeter wavelength range.

3. Using the 43 GHz VLA data, we confirm the predictions of the models made by Anantharamaiah et al. (2000) for RRLs. The integrated $\text{H}53\alpha$ line flux density is a factor of ~ 50 times larger than the integrated $\text{H}92\alpha$ line flux density, in agreement with the predictions of Anantharamaiah et al. (2000). Thus, the $\text{H}53\alpha$ line traces the high-density ($\sim 10^5 \text{ cm}^{-3}$) compact H II regions ($\sim 0.1 \text{ pc}$) in Arp 220.

4. The spatial distributions of the $\text{H}53\alpha$ and the $\text{H}92\alpha$ line emission are similar in the direction of the E component. On the W component the low-density ($\sim 10^3 \text{ cm}^{-3}$) ionized gas component is more extended than the high-density ($\sim 10^5 \text{ cm}^{-3}$) component.

5. The kinematic and spatial distribution behavior as observed in the RRL $\text{H}53\alpha$ is in agreement with results reported from CO and H I observations (Sakamoto et al. 1999; Mundell et al. 2001), supporting the counterrotation of two disks in Arp 220 at smaller scales.

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